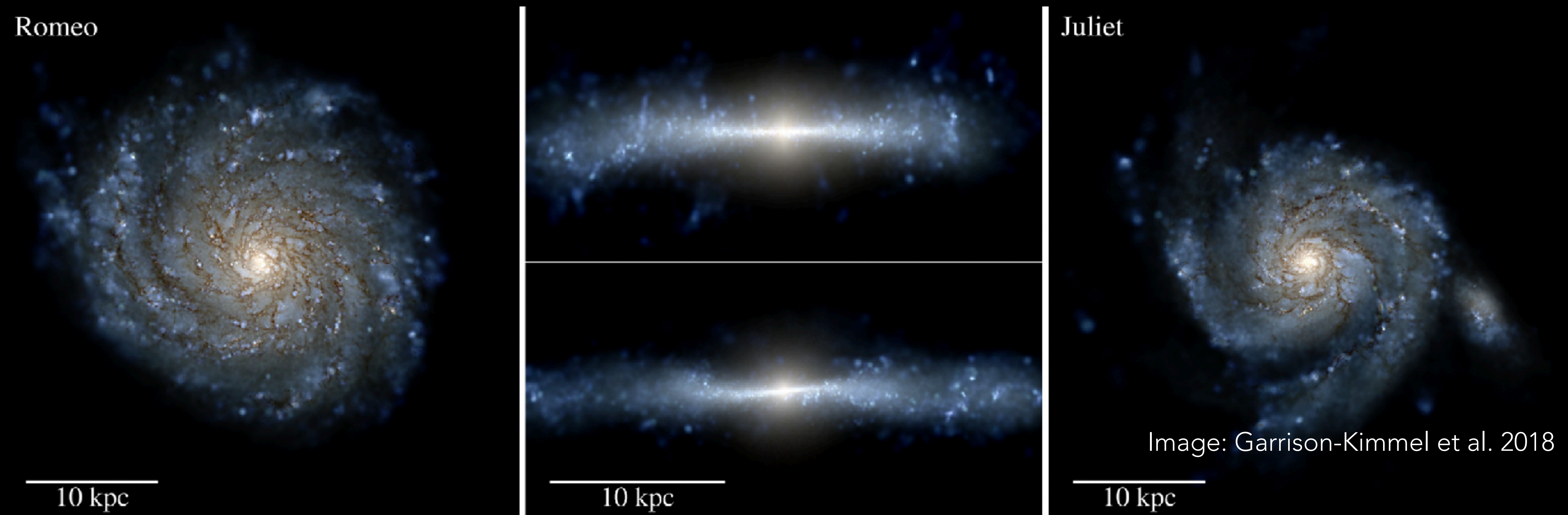
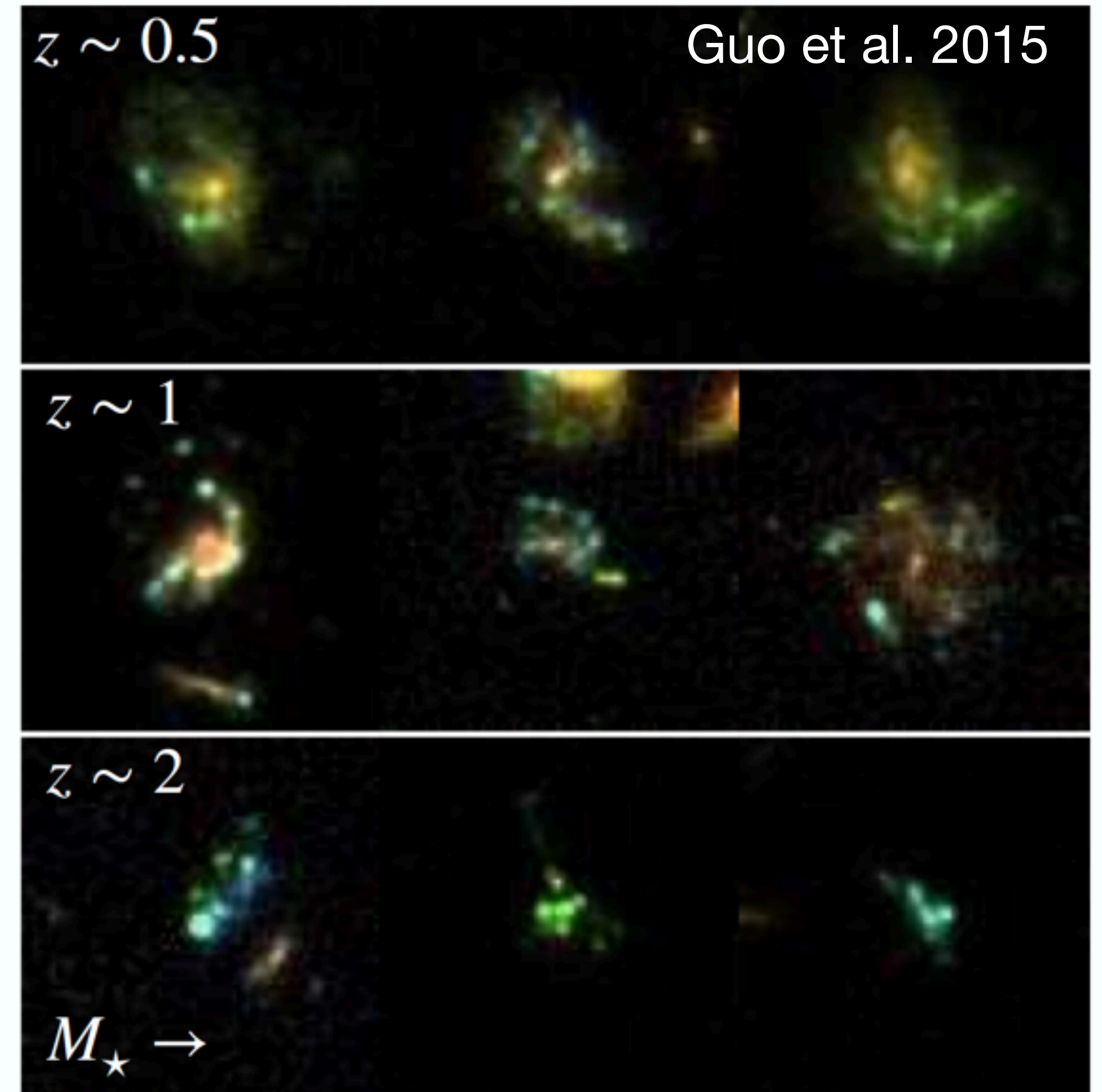
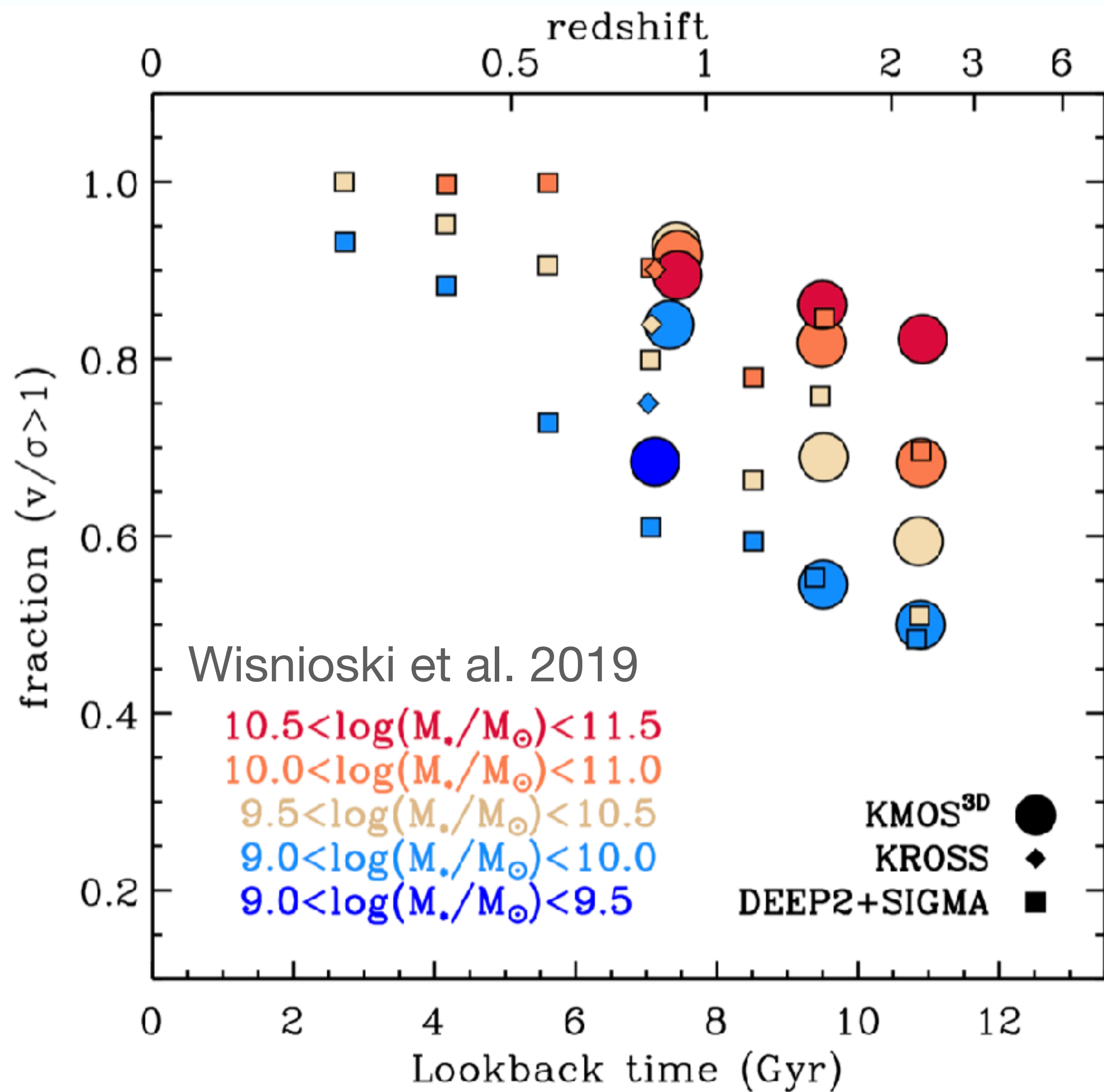


THE PHYSICS OF GALAXY DISK FORMATION



James Bullock (UC Irvine)

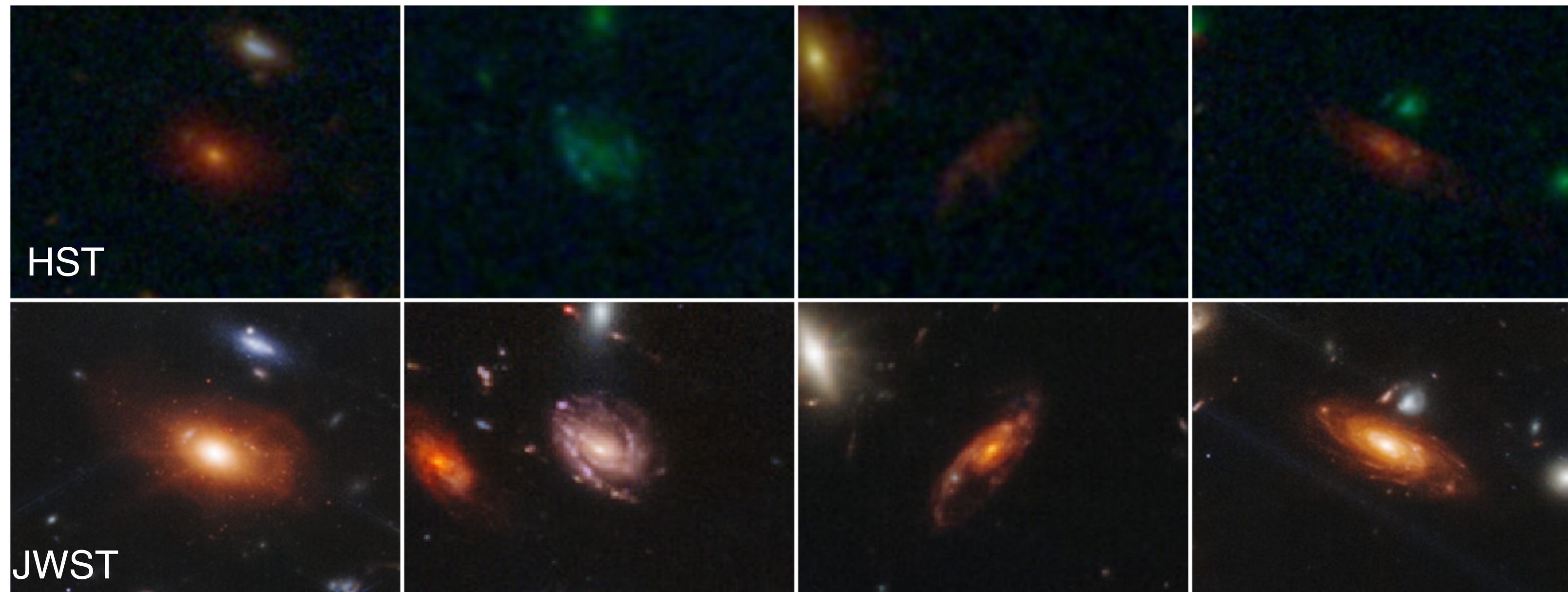


e.g. Elmegreen & Elmegreen 05; Oesch +10;
Guo+15, Margalef-Bentabol +18; Whitney +21

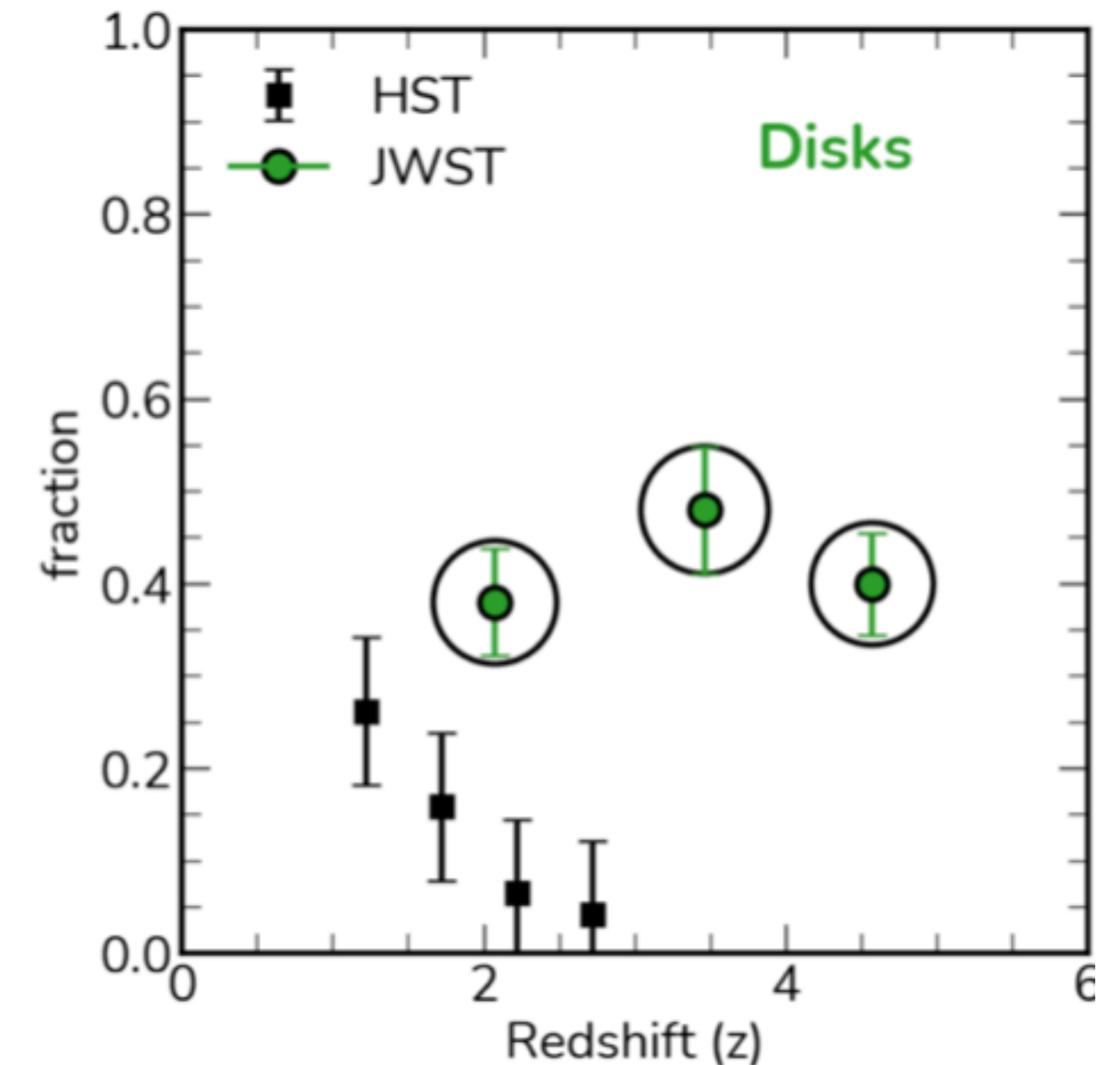
Among star-forming galaxies:
More disks @ high mass & low z

High-z disks:
irregular & clumpy

JWST sees more high-z disk galaxies than expected
->Current ground-based IFS may underestimate V/σ ?



Ferreira et al. 2022



Also: Robertson et al 2022, Kartaltepe et al. 2022, Colin et al. 2022, Vega-Ferrero et al. 2023, ...

Bold claim: Hard to say what it means... We really don't have a fully agreed-upon "physics-based" theory for disk formation.

Towards realistic disks in LCDM simulations

THE ASTROPHYSICAL JOURNAL, 742:76 (10pp), 2011 December 1
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2011

doi:10.1088/0004-637X/742/2/76

FORMING REALISTIC LATE-TYPE SPIRALS IN A Λ CDM UNIVERSE: THE ERIS SIMULATION

JAVIERA GUEDES^{1,2}, SIMONE CALLEGARI³, PIERO MADAU¹, AND LUCIO MAYER^{2,3}

THE ASTROPHYSICAL JOURNAL, 773:43 (19pp), 2013 August 10
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2013

doi:10.1088/0004-637X/773/1/43

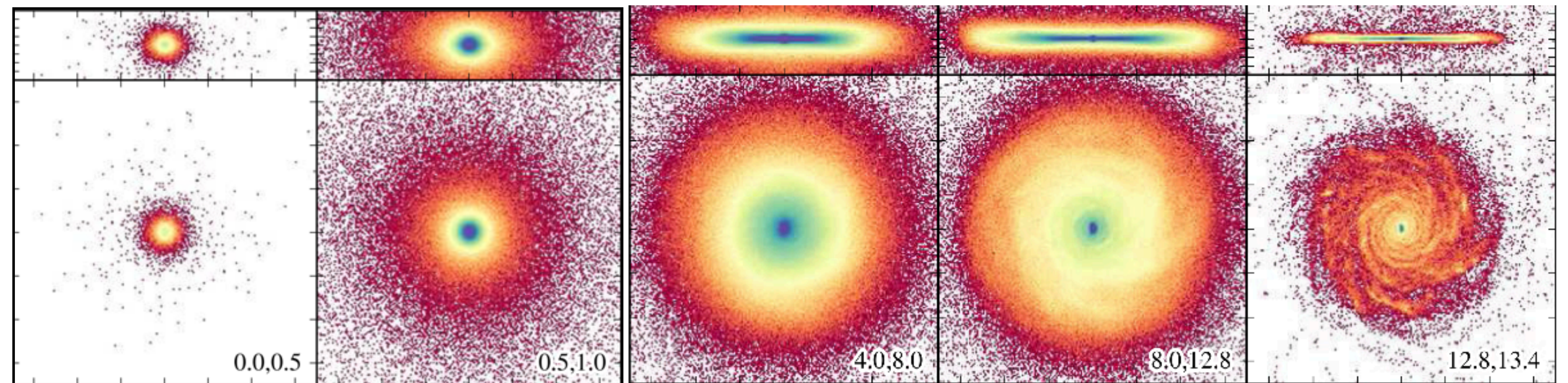
INSIDE OUT AND UPSIDE DOWN: TRACING THE ASSEMBLY OF A SIMULATED DISK GALAXY USING MONO-AGE STELLAR POPULATIONS

JONATHAN C. BIRD^{1,2,3,8}, STELIOS KAZANTZIDIS^{1,2}, DAVID H. WEINBERG^{1,2}, JAVIERA GUEDES⁴,
SIMONE CALLEGARI⁵, LUCIO MAYER⁶, AND PIERO MADAU⁷

$z \sim 10$

$z \sim 0.1$

$z = 0$



Cosmic time



15 kpc

(See also Thacker & Couchman 2000; Governato et al. 2004; Robertson et al. 2006; Giuseppe et al. 2015...)

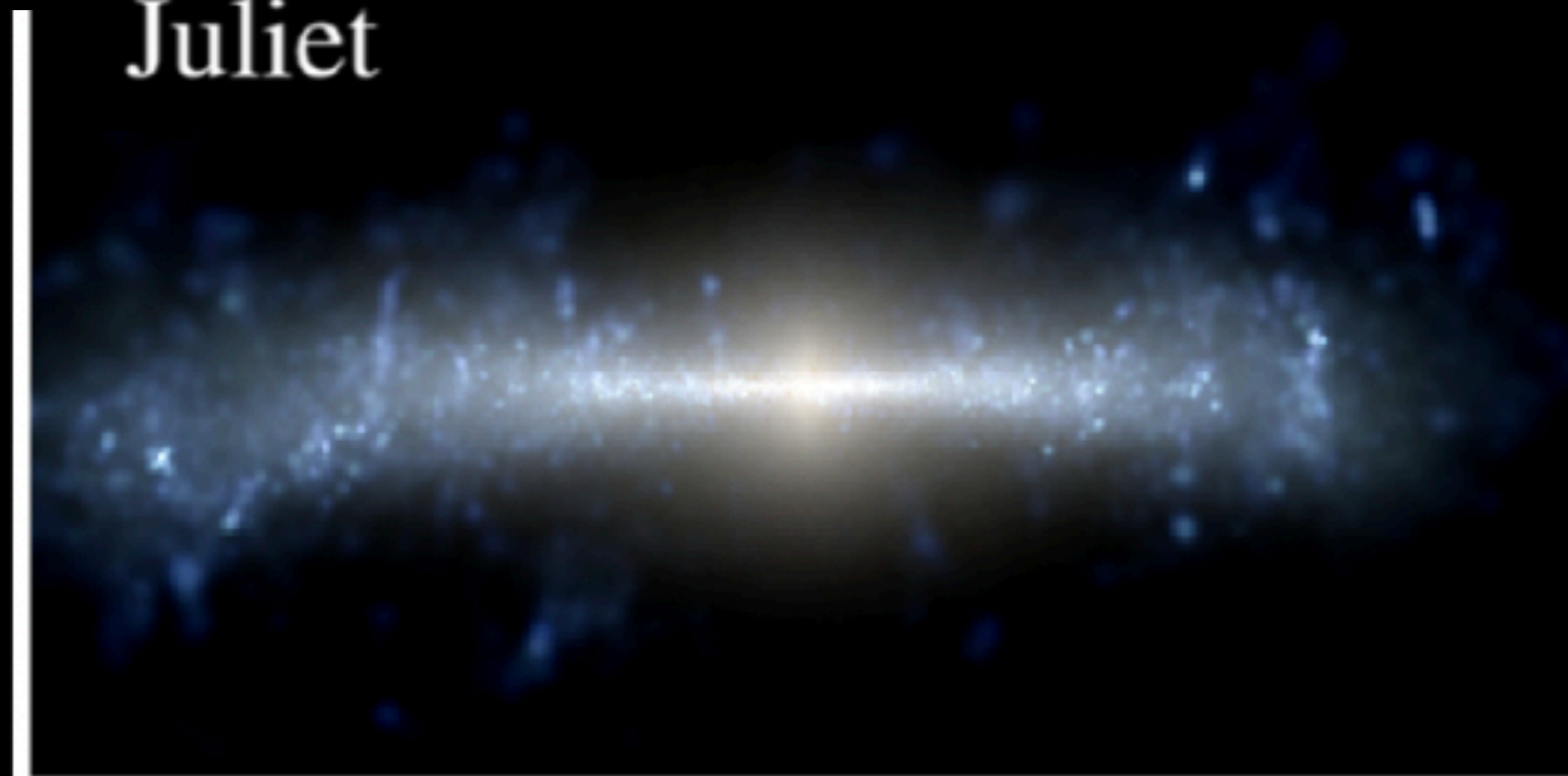


Simulations of Milky-Way size galaxies

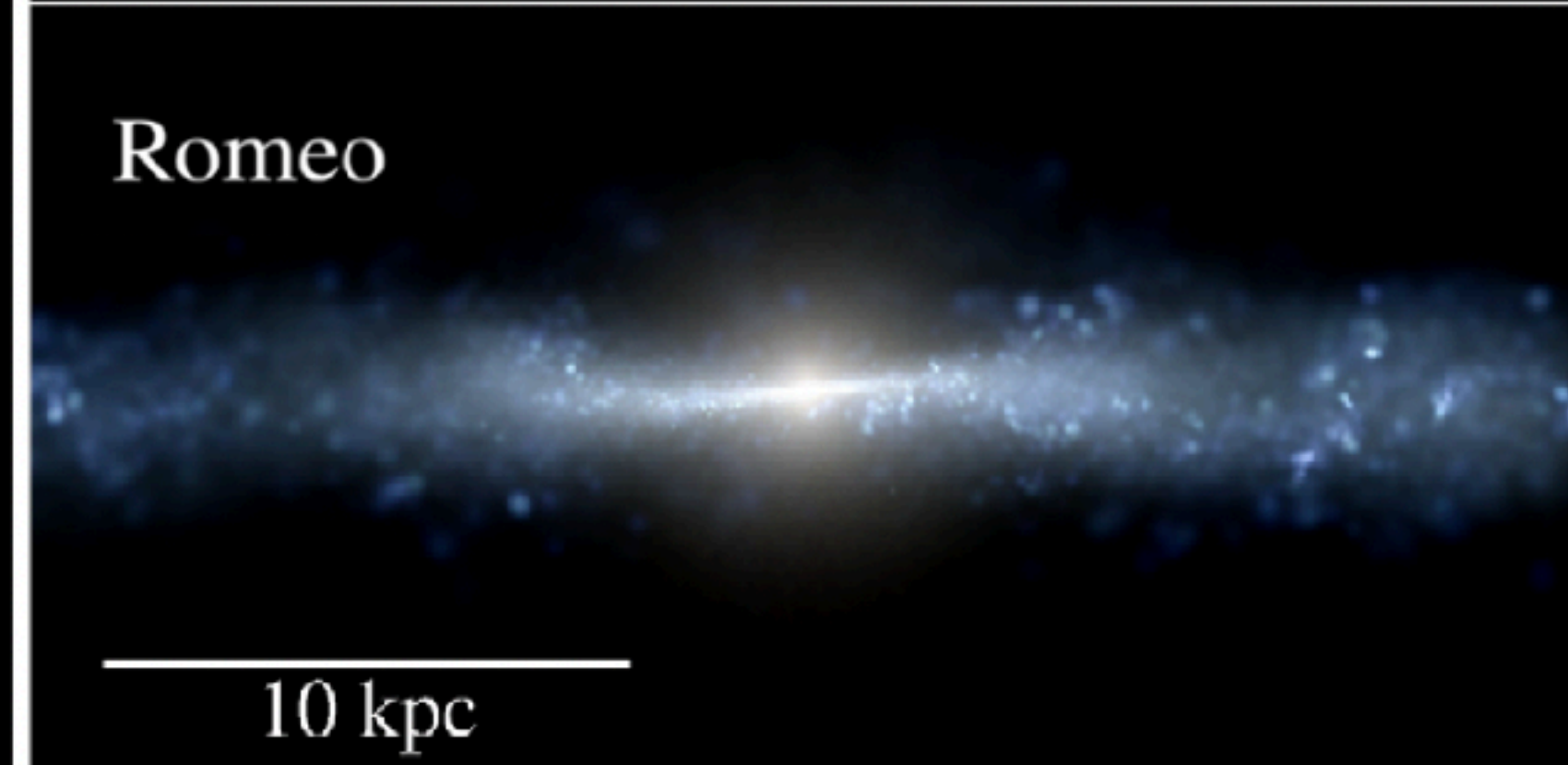
Romeo



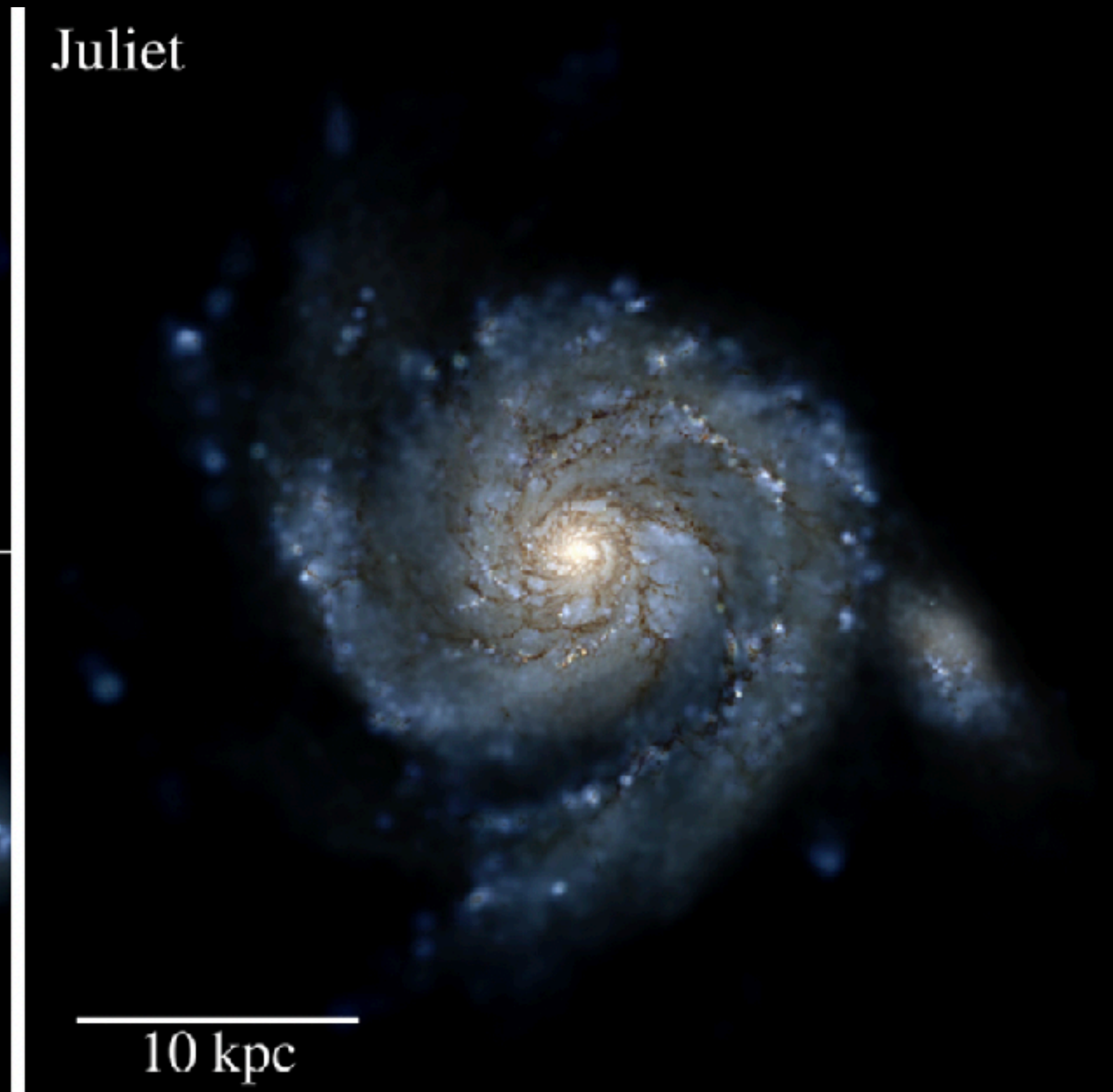
Juliet



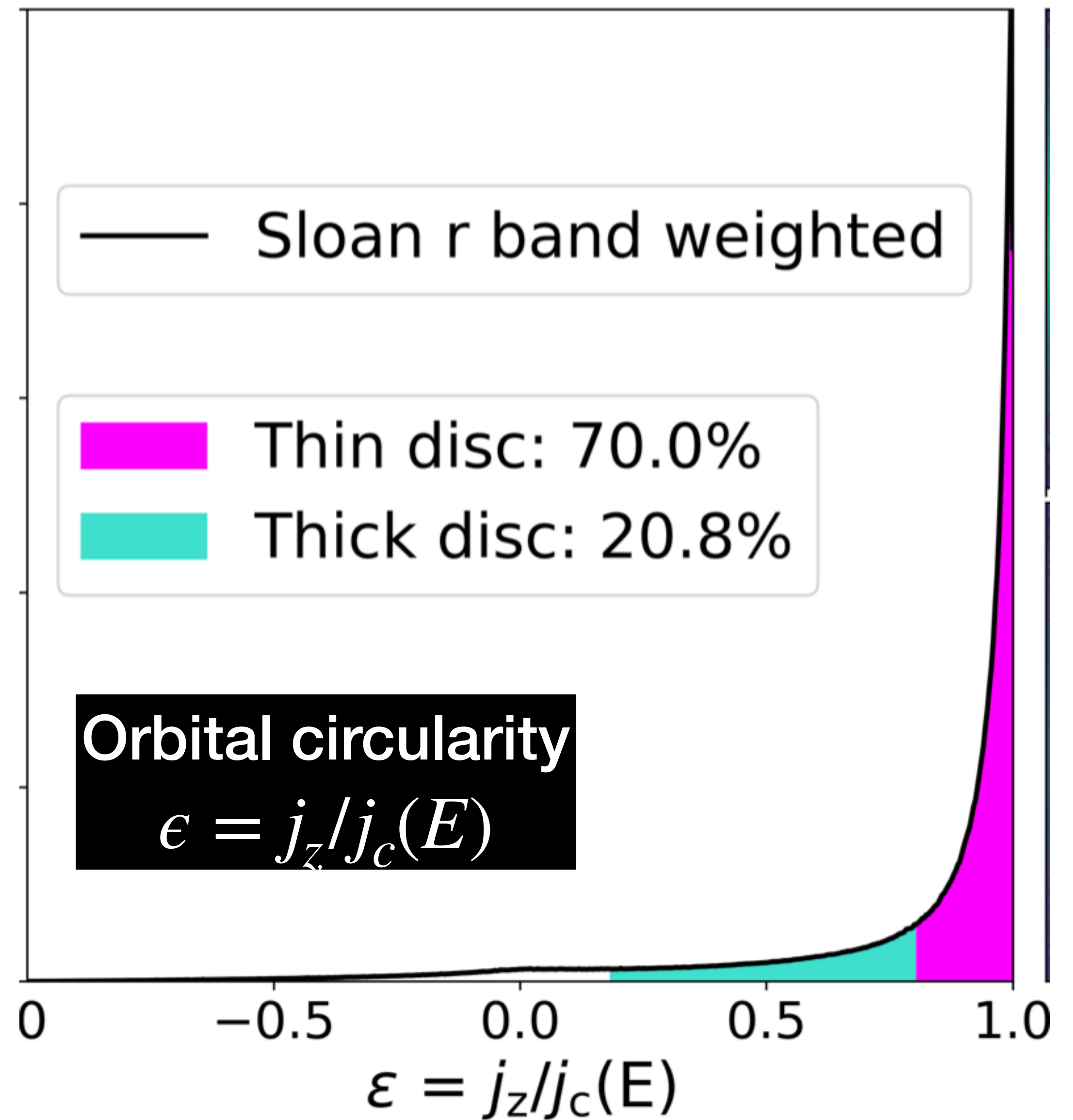
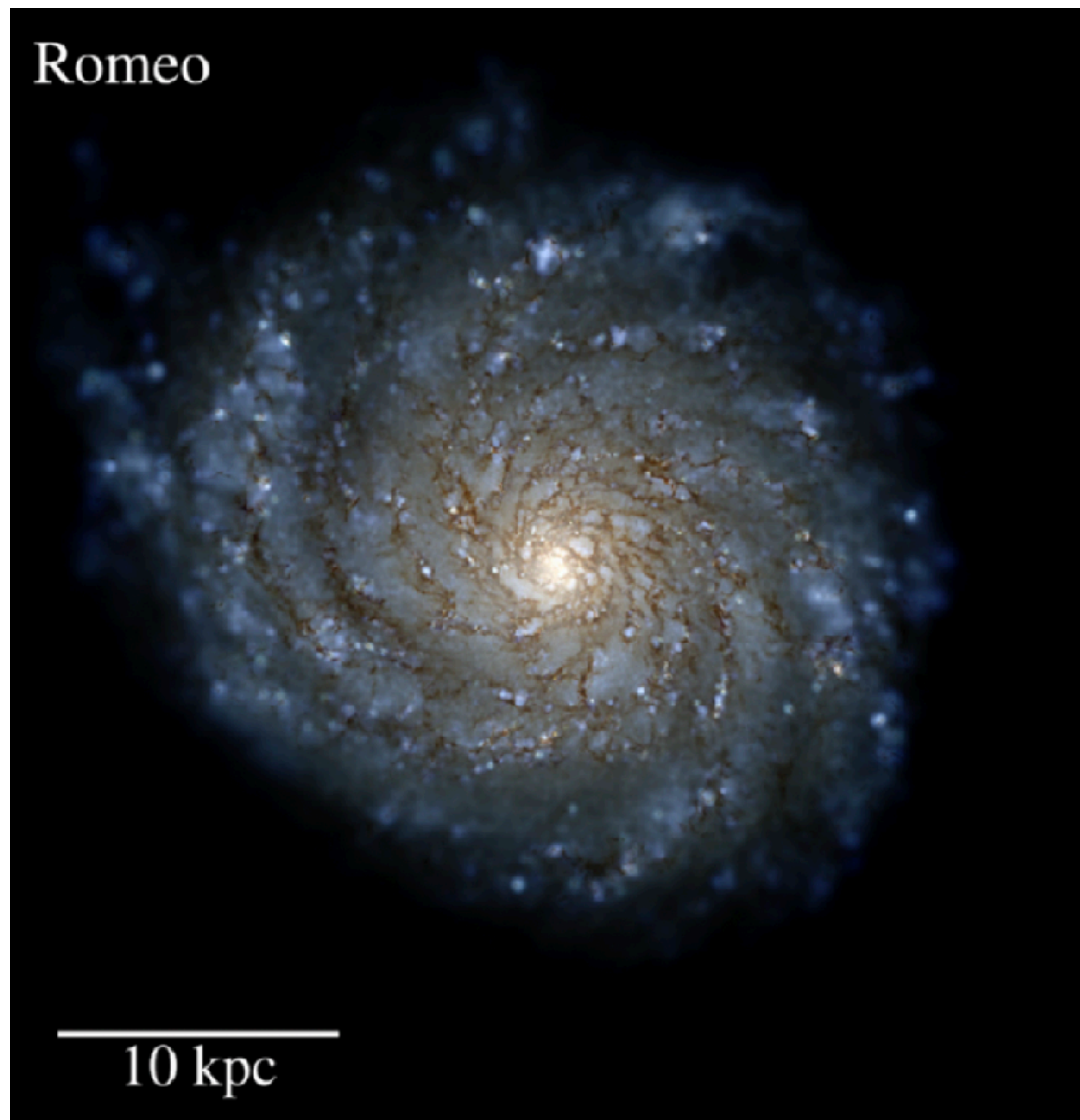
Romeo



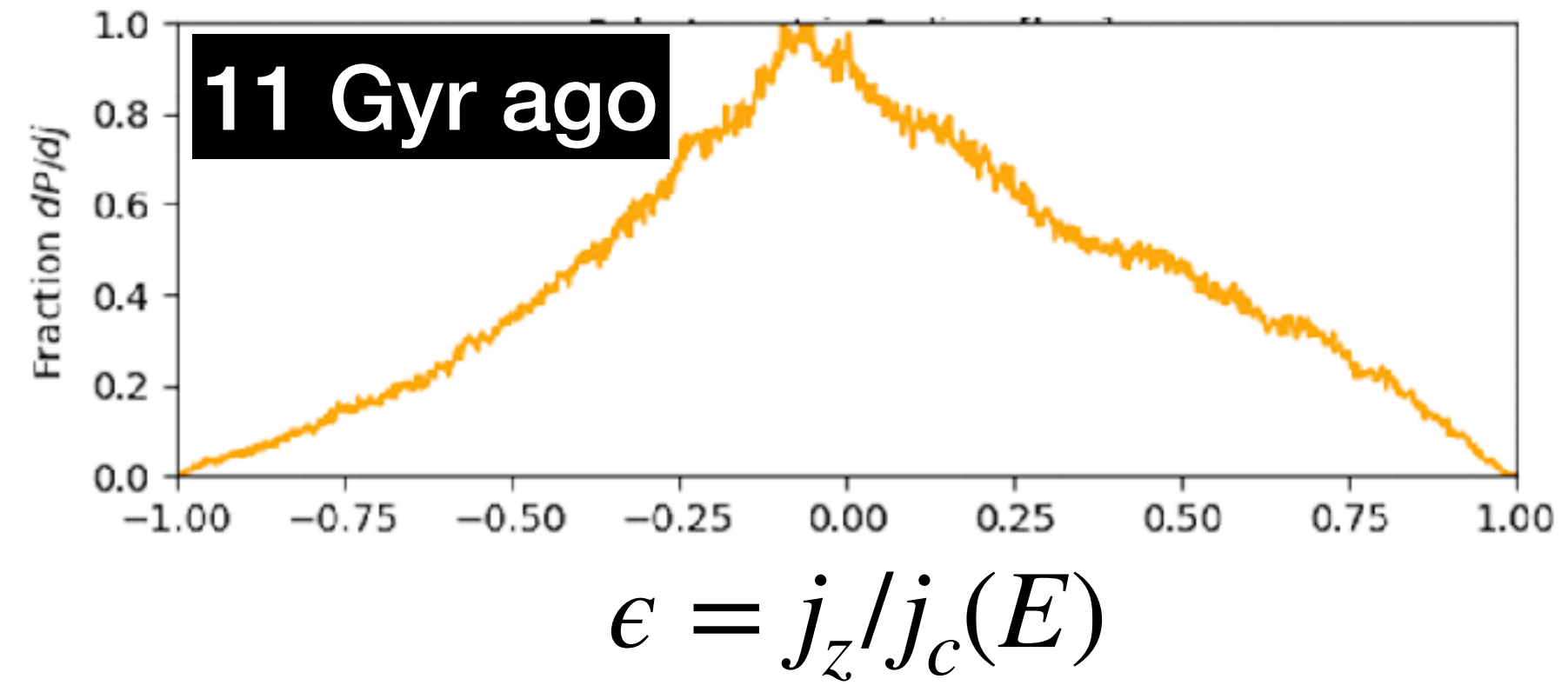
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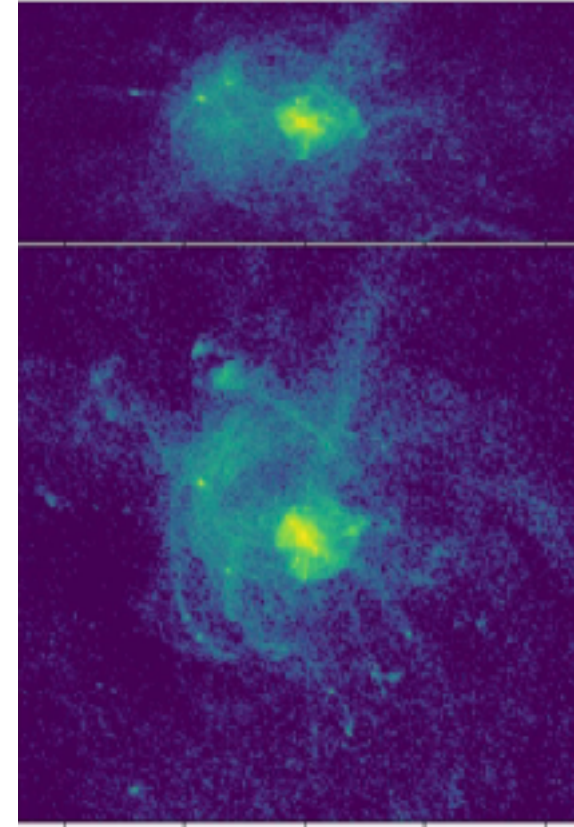
This disk is very thin at $z=0$



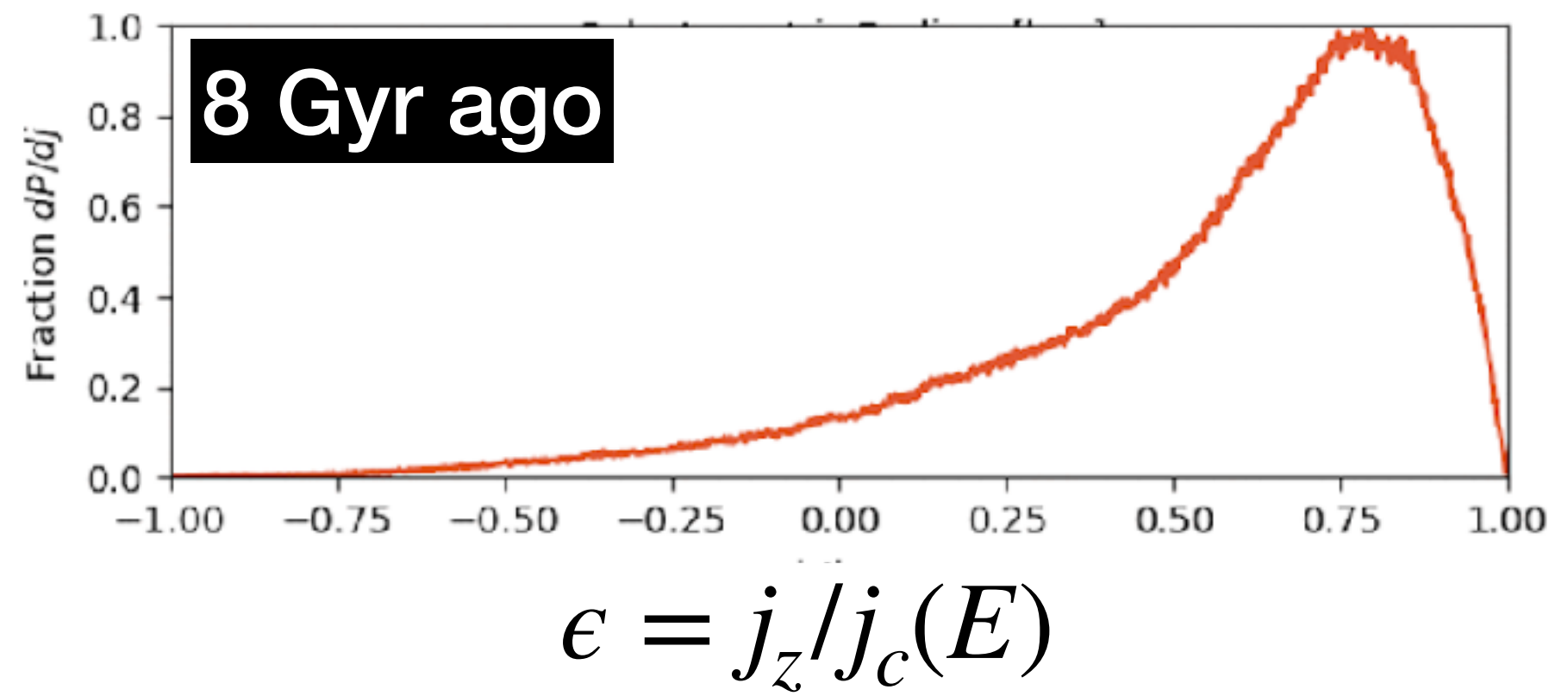
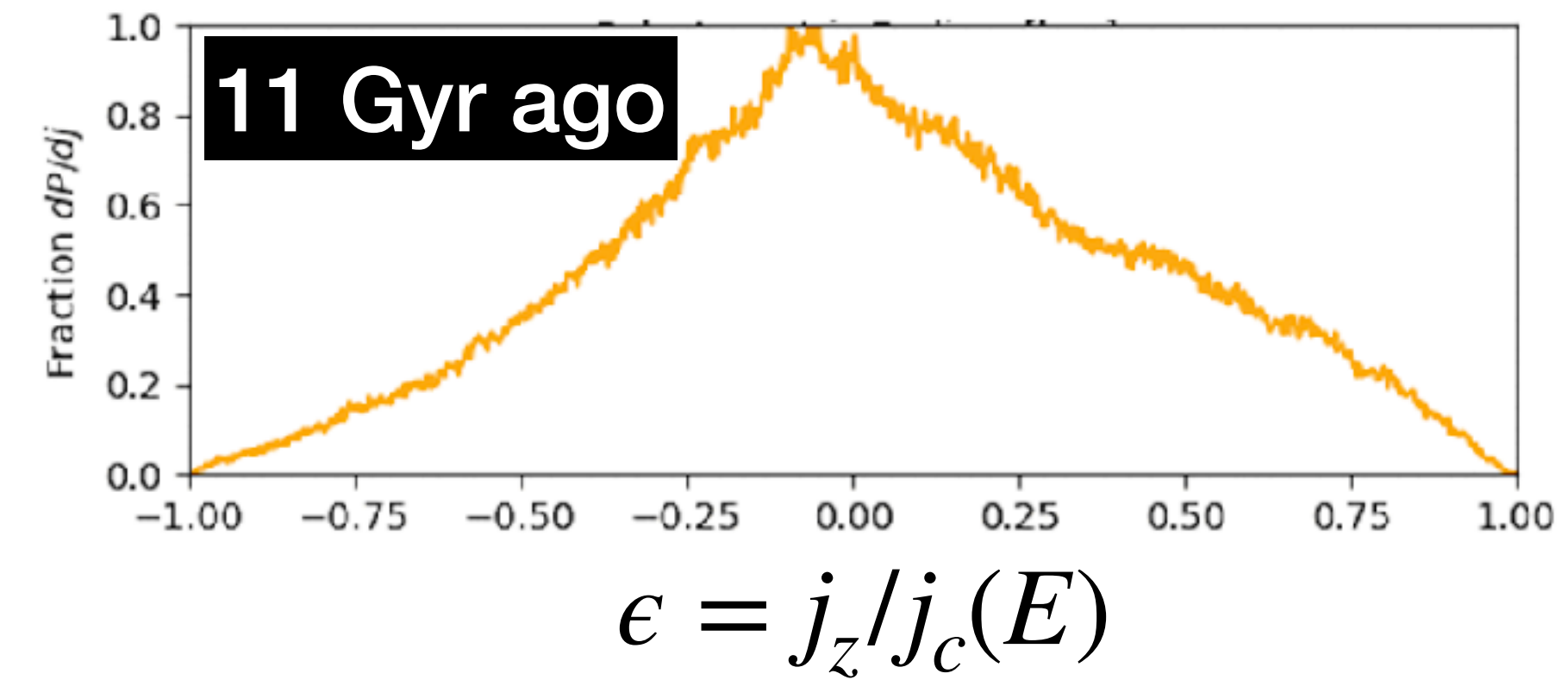
Orbits of young stars



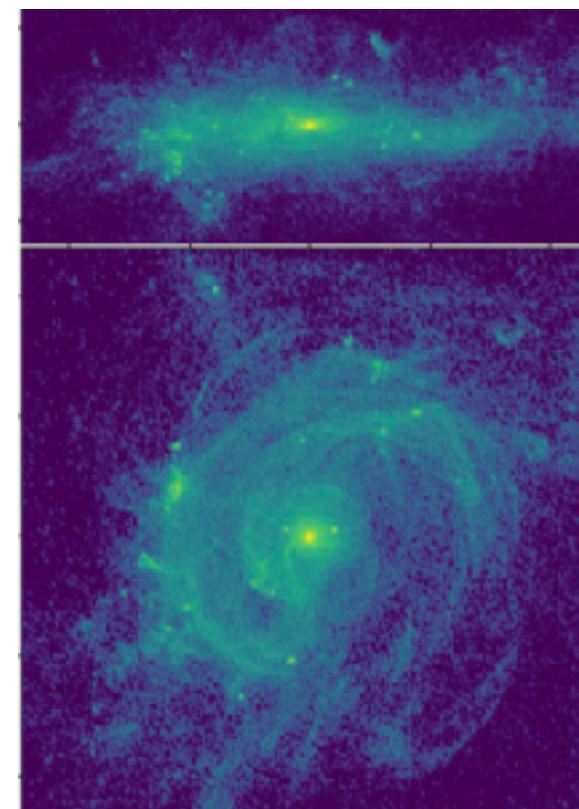
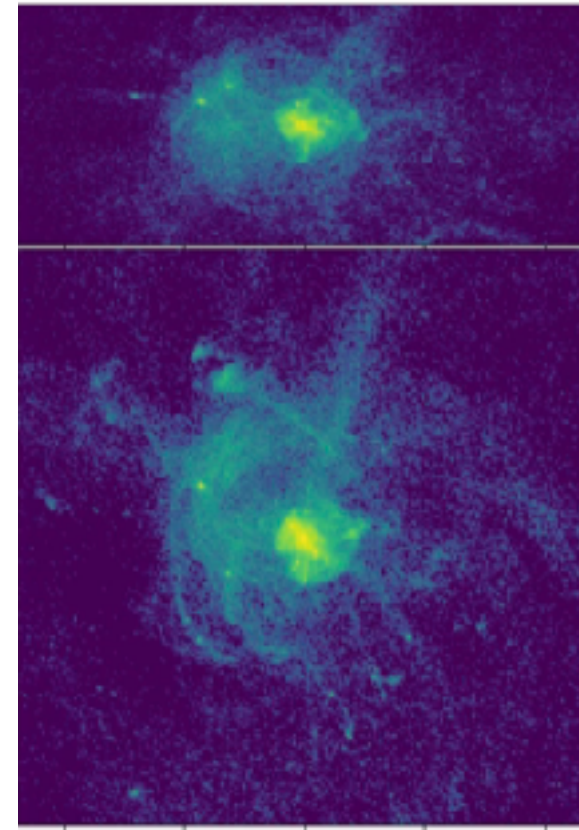
Young stars



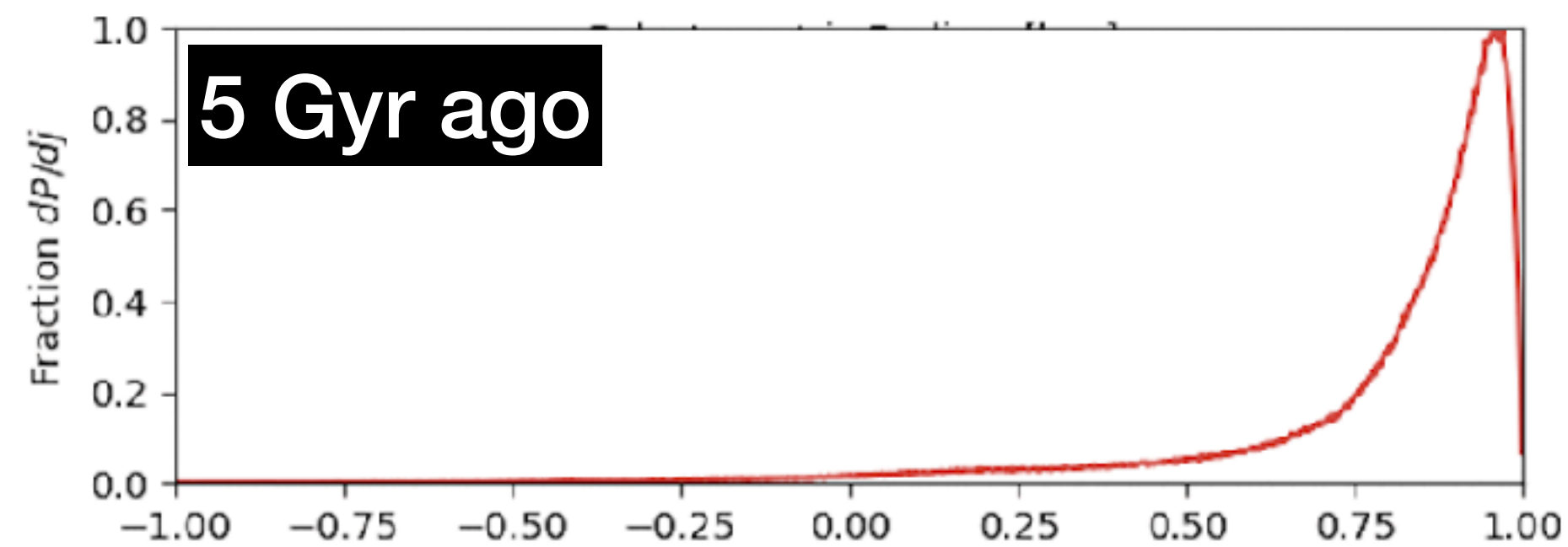
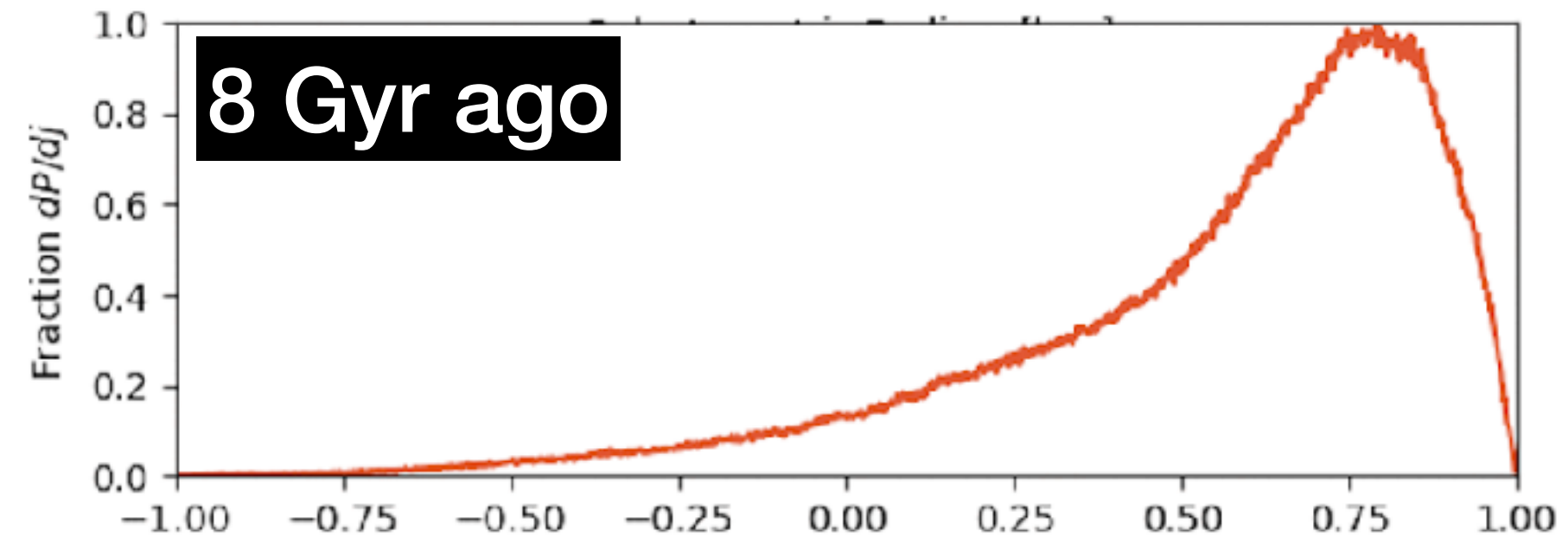
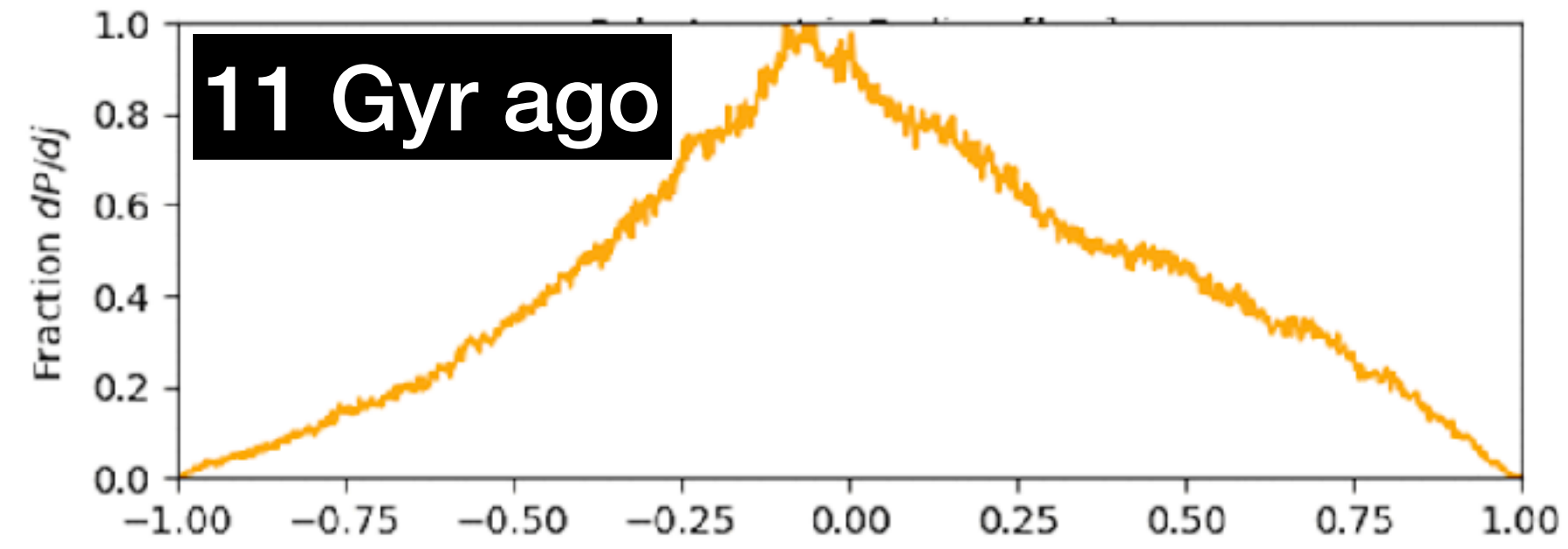
Orbits of young stars



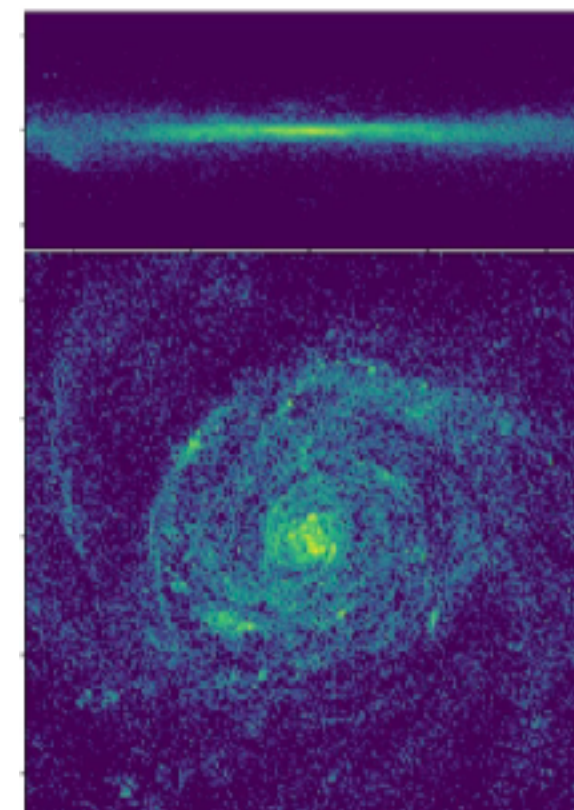
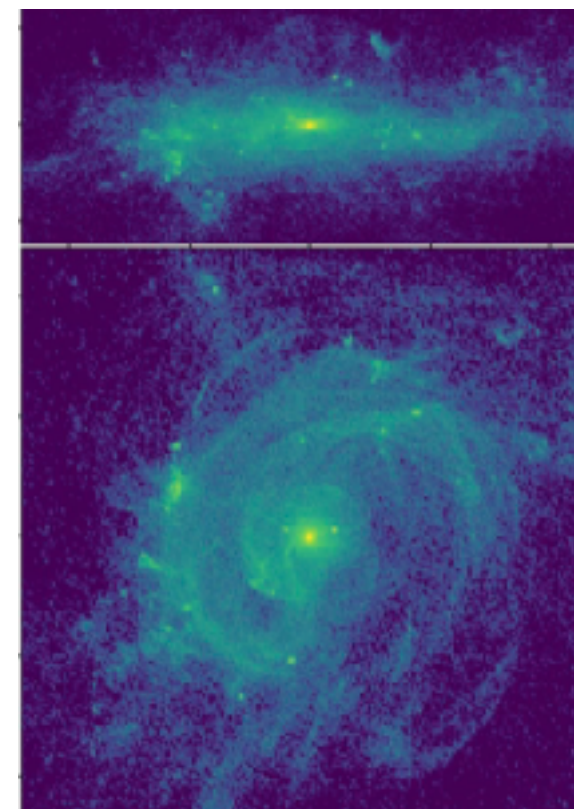
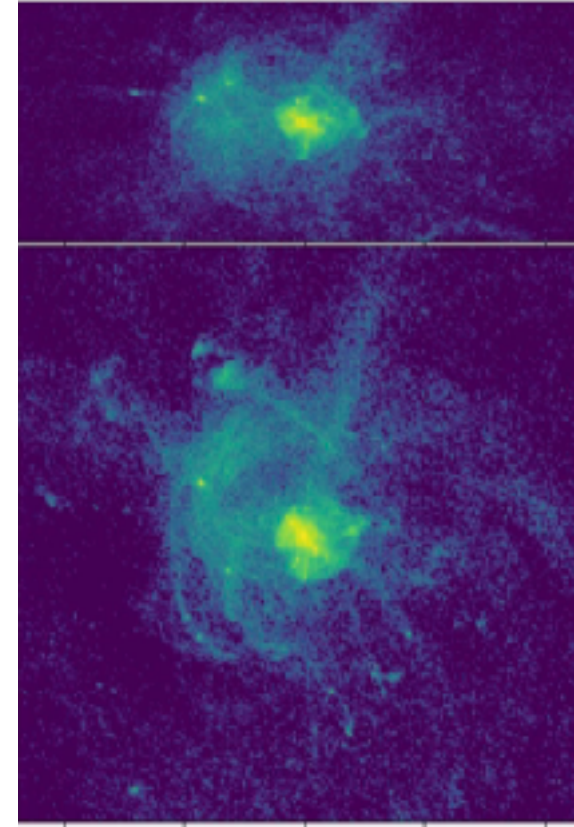
Young stars



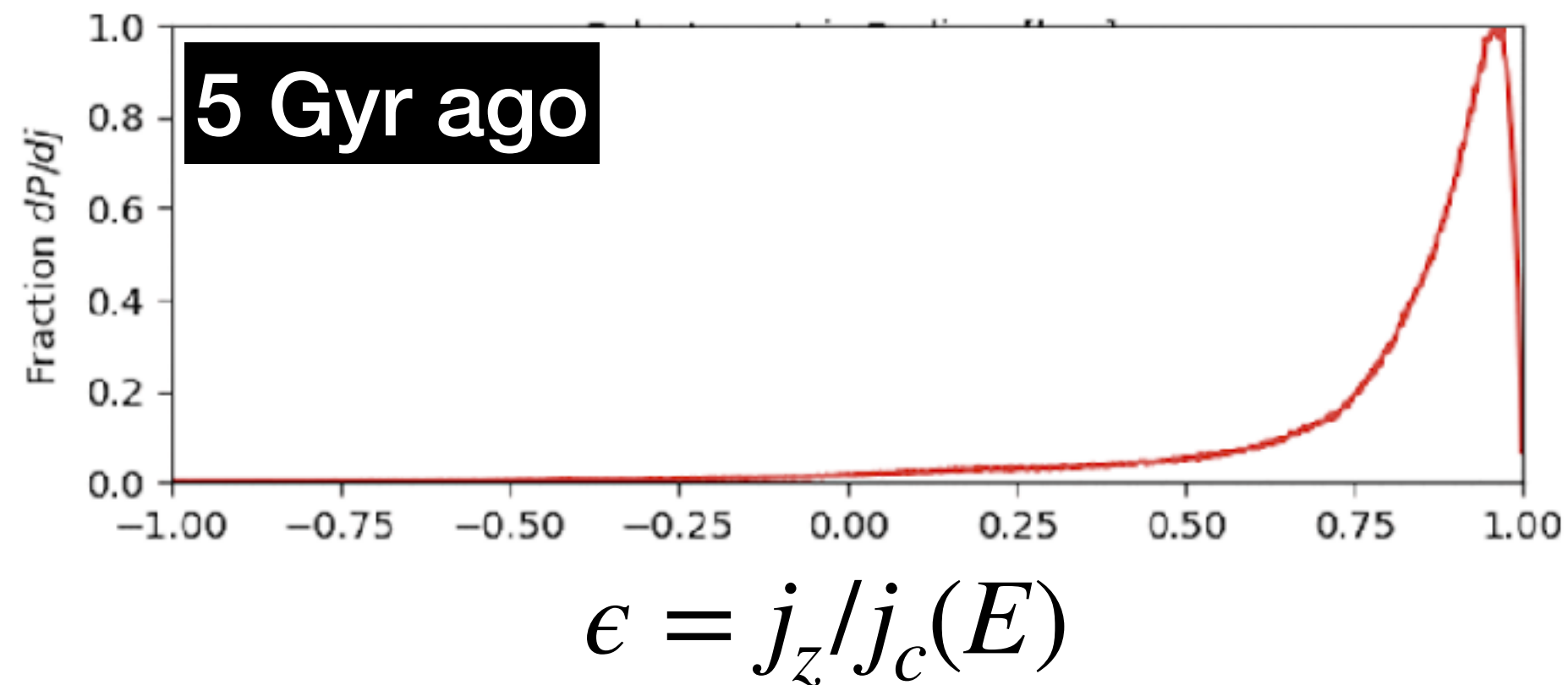
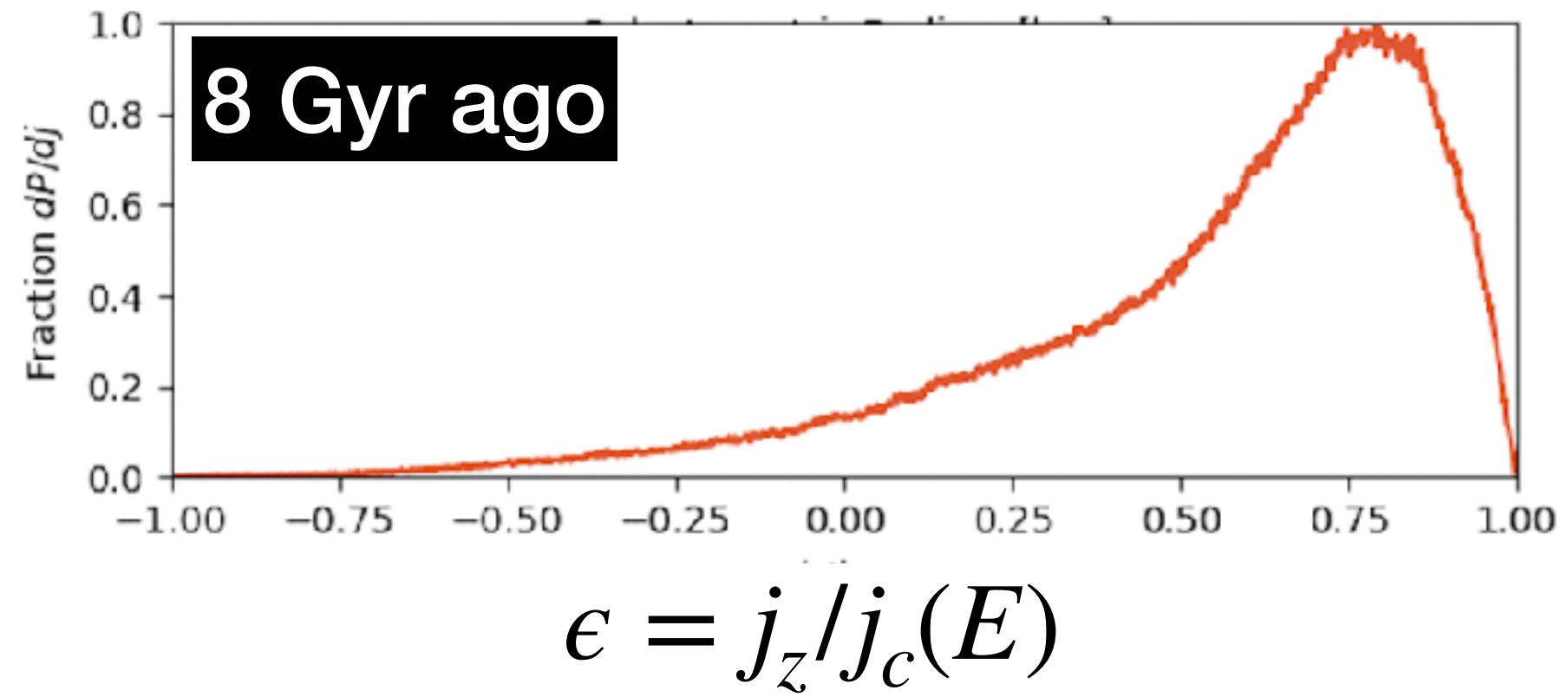
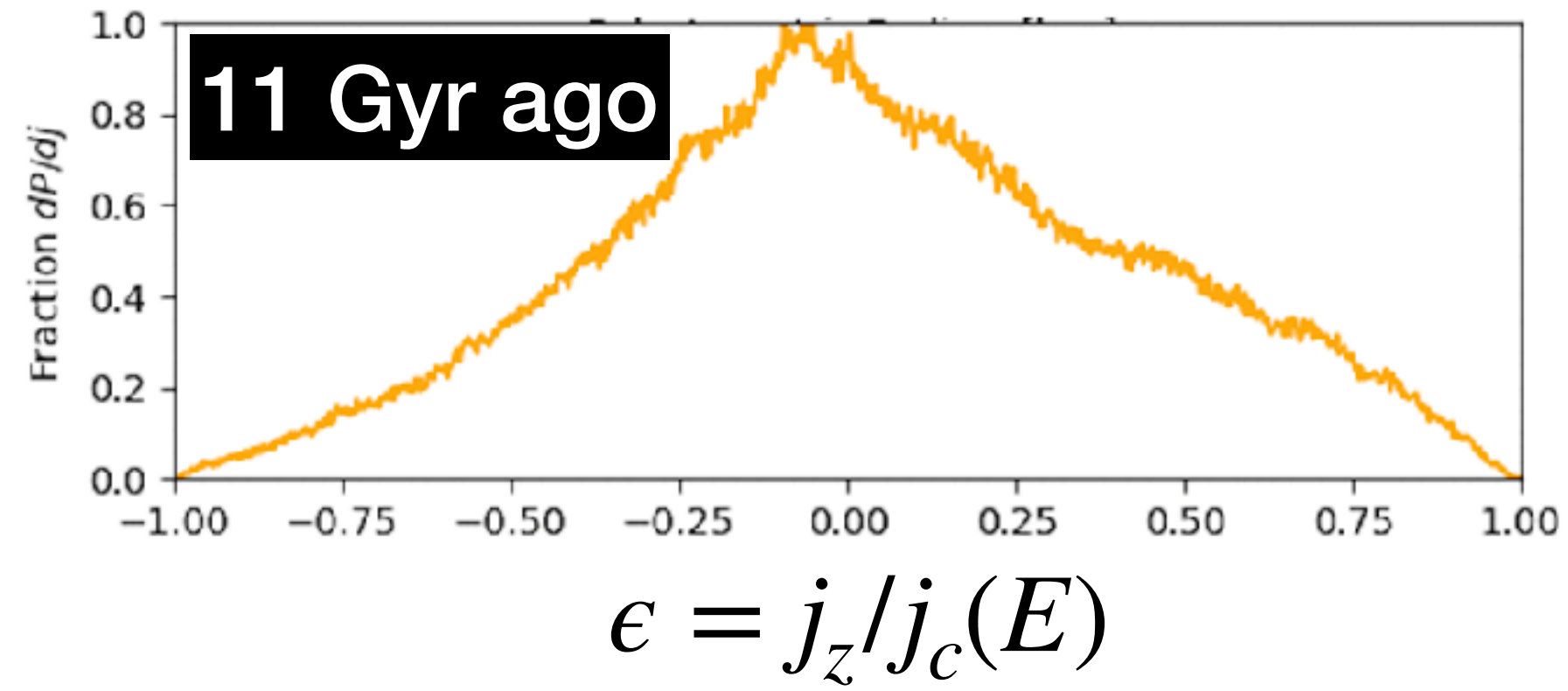
Orbits of young stars



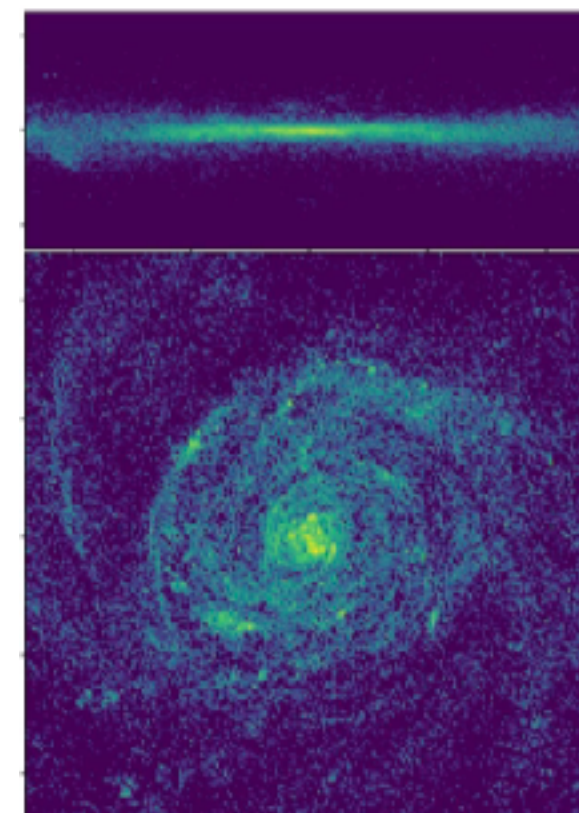
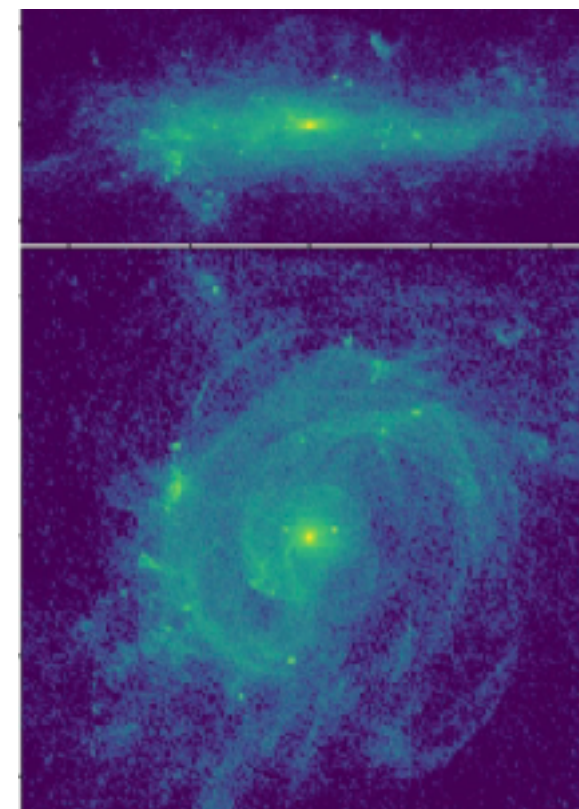
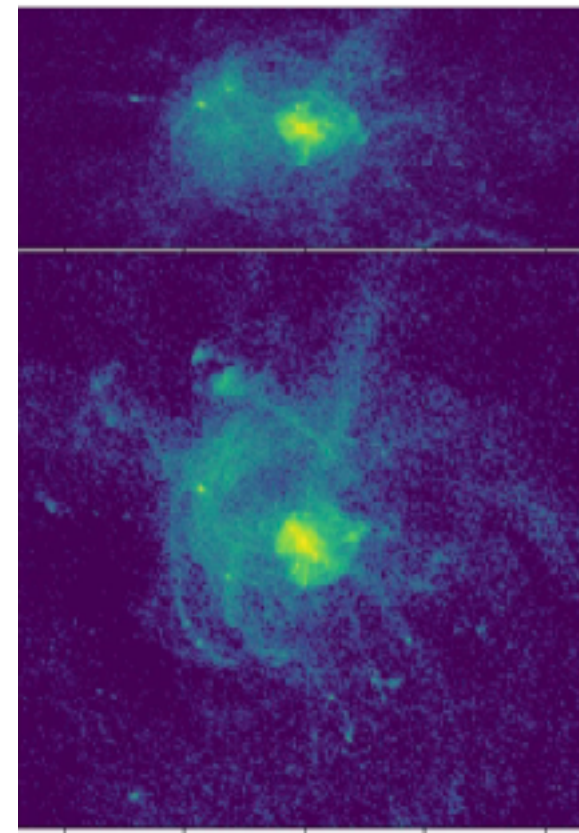
Young stars



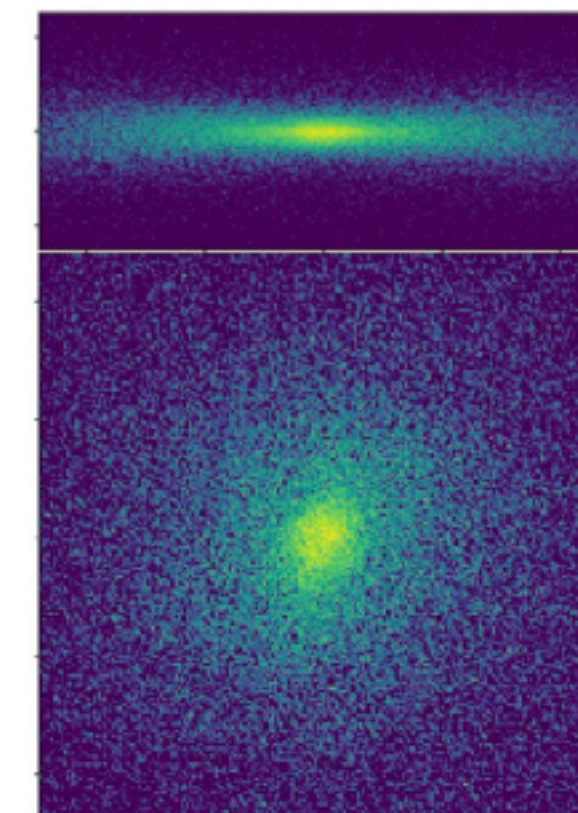
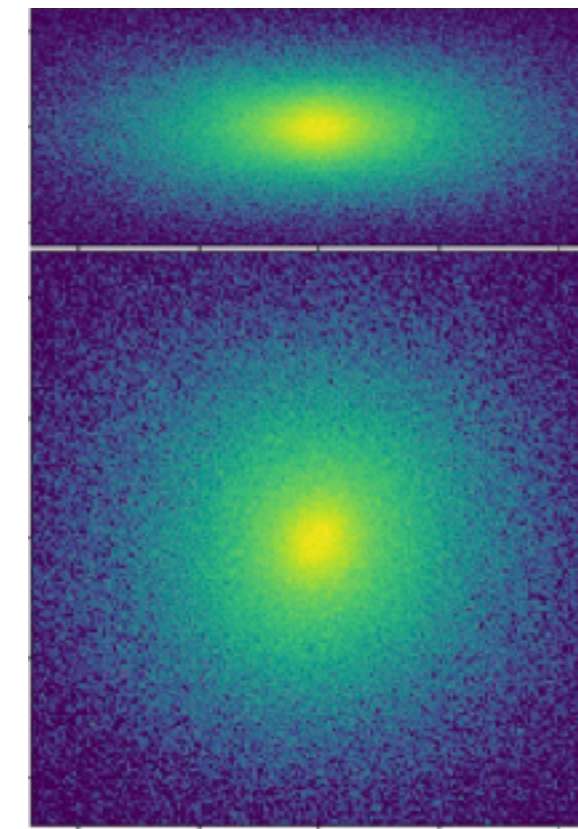
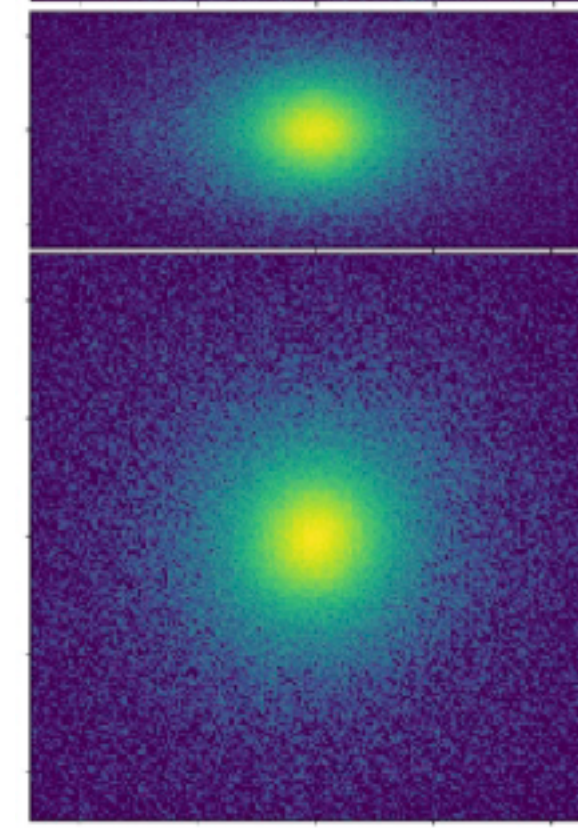
Orbits of young stars



Young stars



Same stars at $z=0$



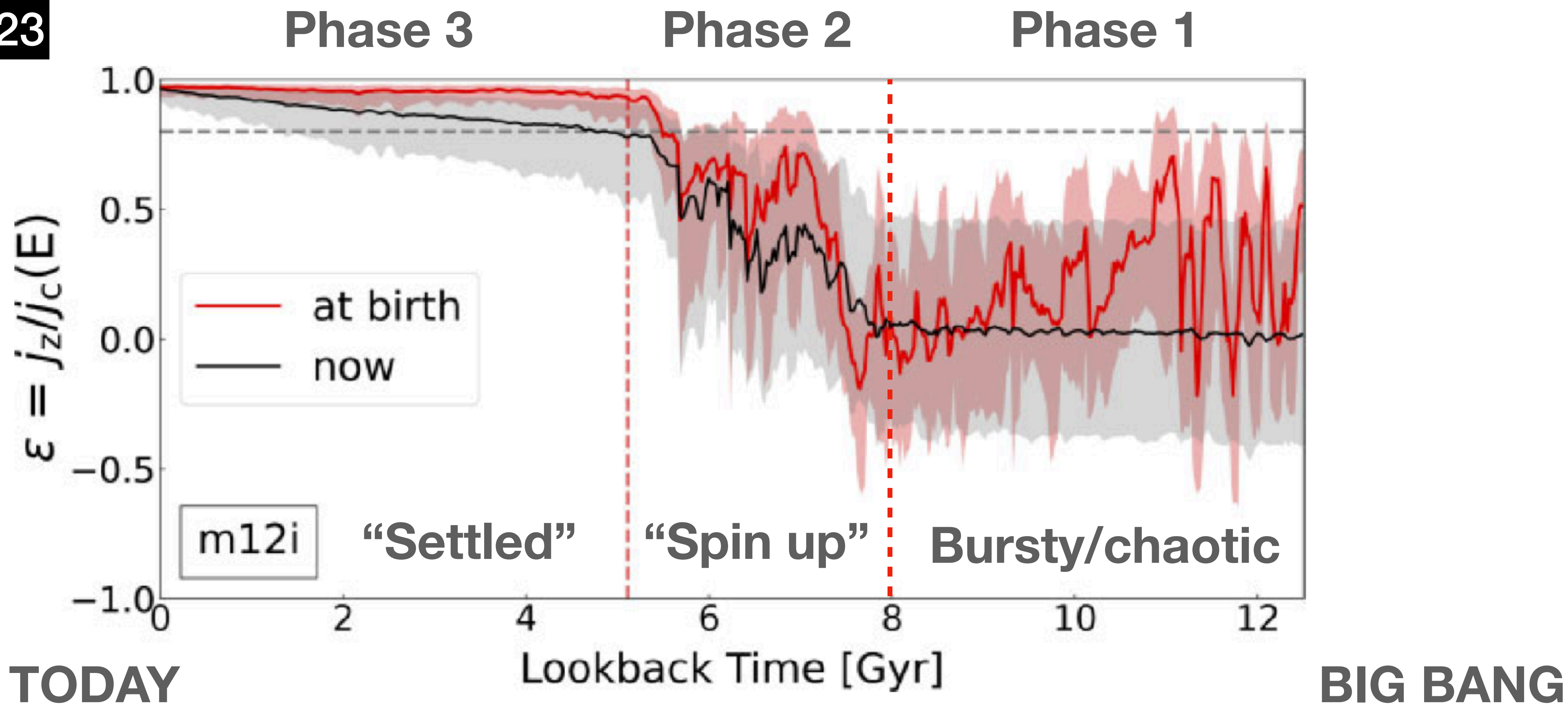
“Spheroid”

“Thick disk”

“Thin disk”

3 Phases of Galaxy Disk Formation

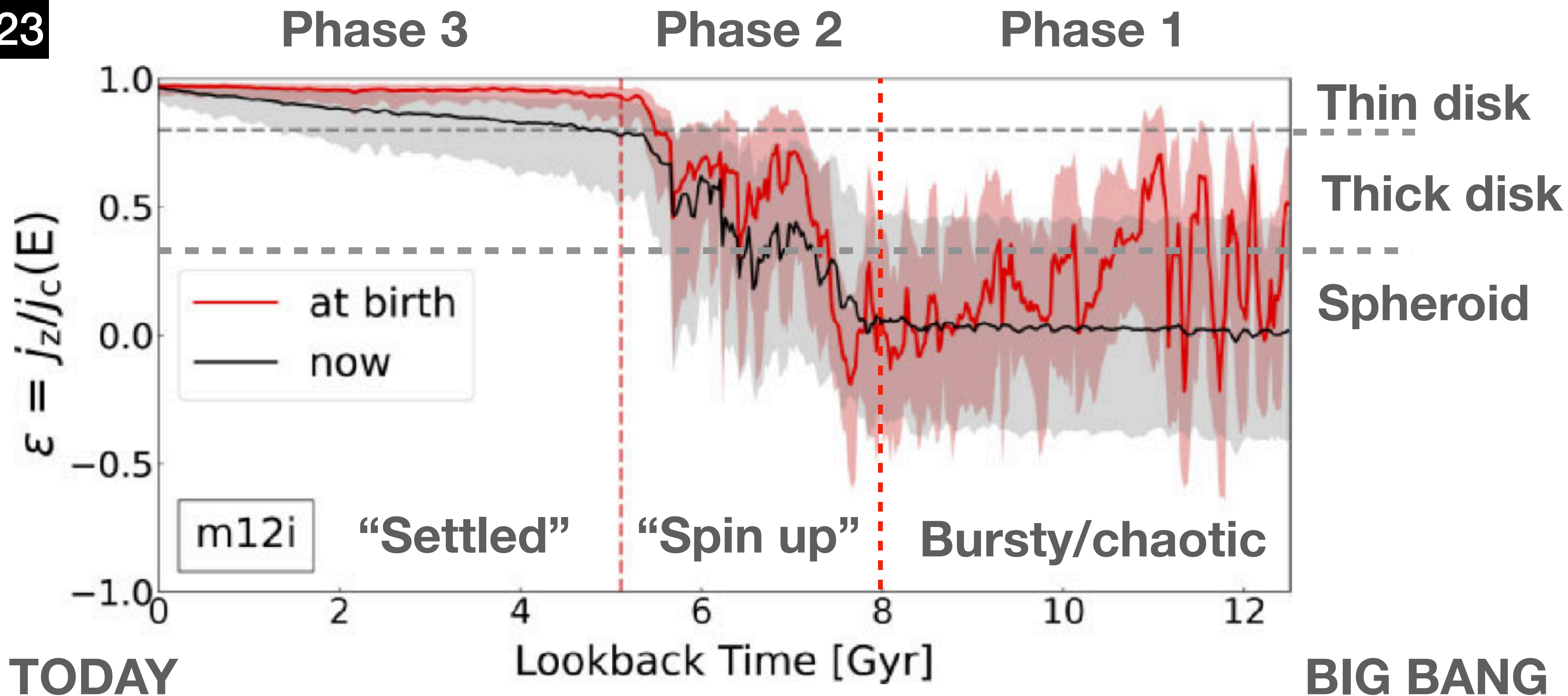
Yu+2023



See also Fiona McCluskey et al. 2023

3 Phases of Galaxy Disk Formation

Yu+2023



See also Fiona McCluskey et al. 2023

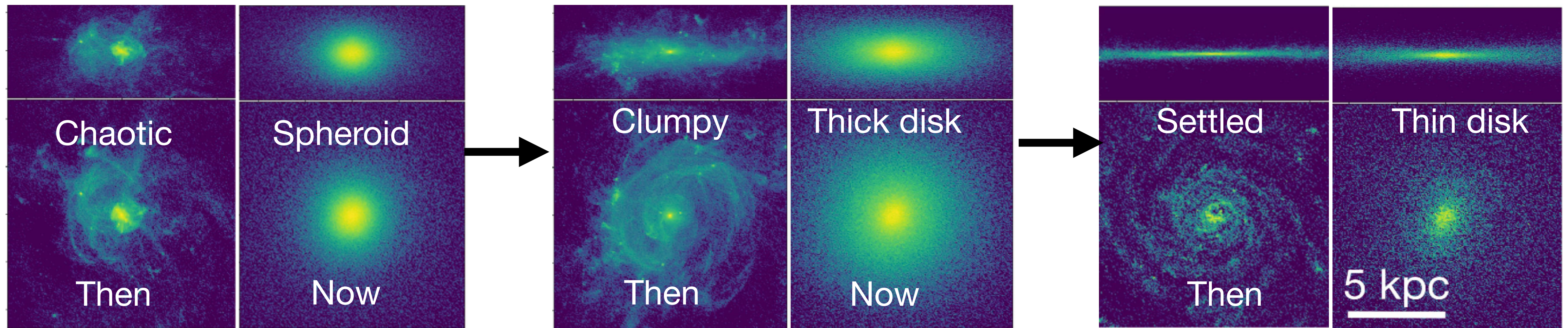
3 Phases of Galaxy Disk Formation

Yu+2023

Phase 1

Phase 2

Phase 3



Young stars

Same stars
 $z=0$

Young stars

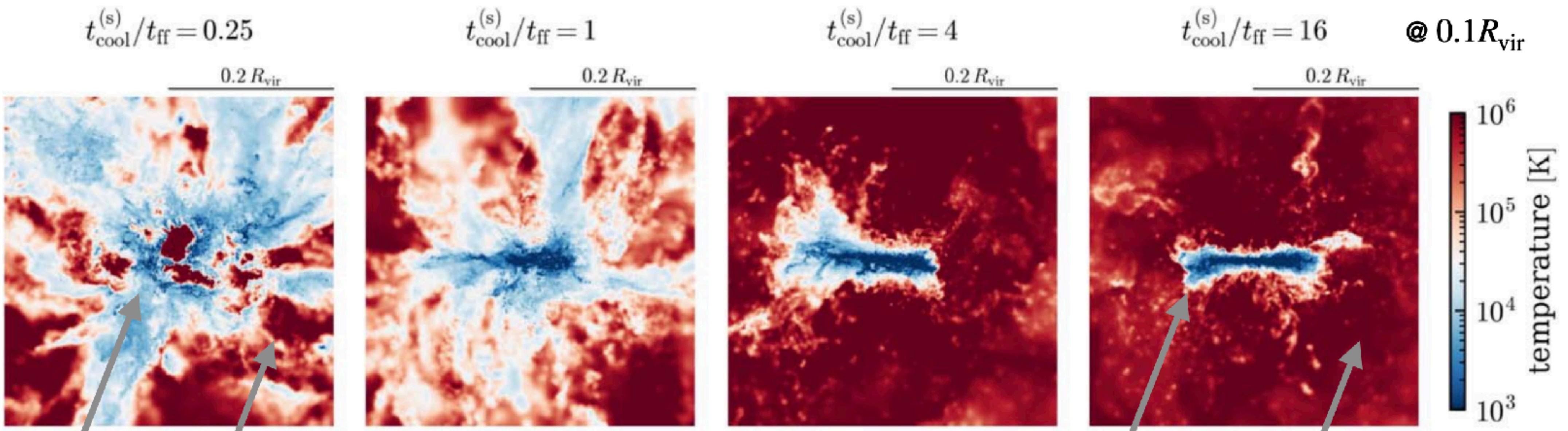
Same stars
 $z=0$

Young stars

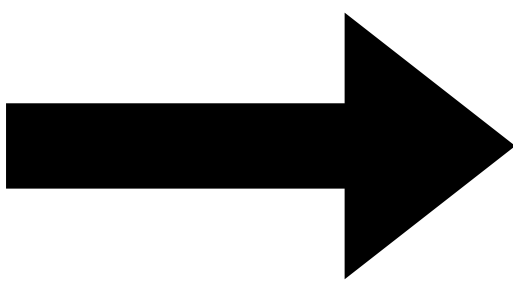
Same stars
 $z=0$

Inner CGM virialization in *FIRE*

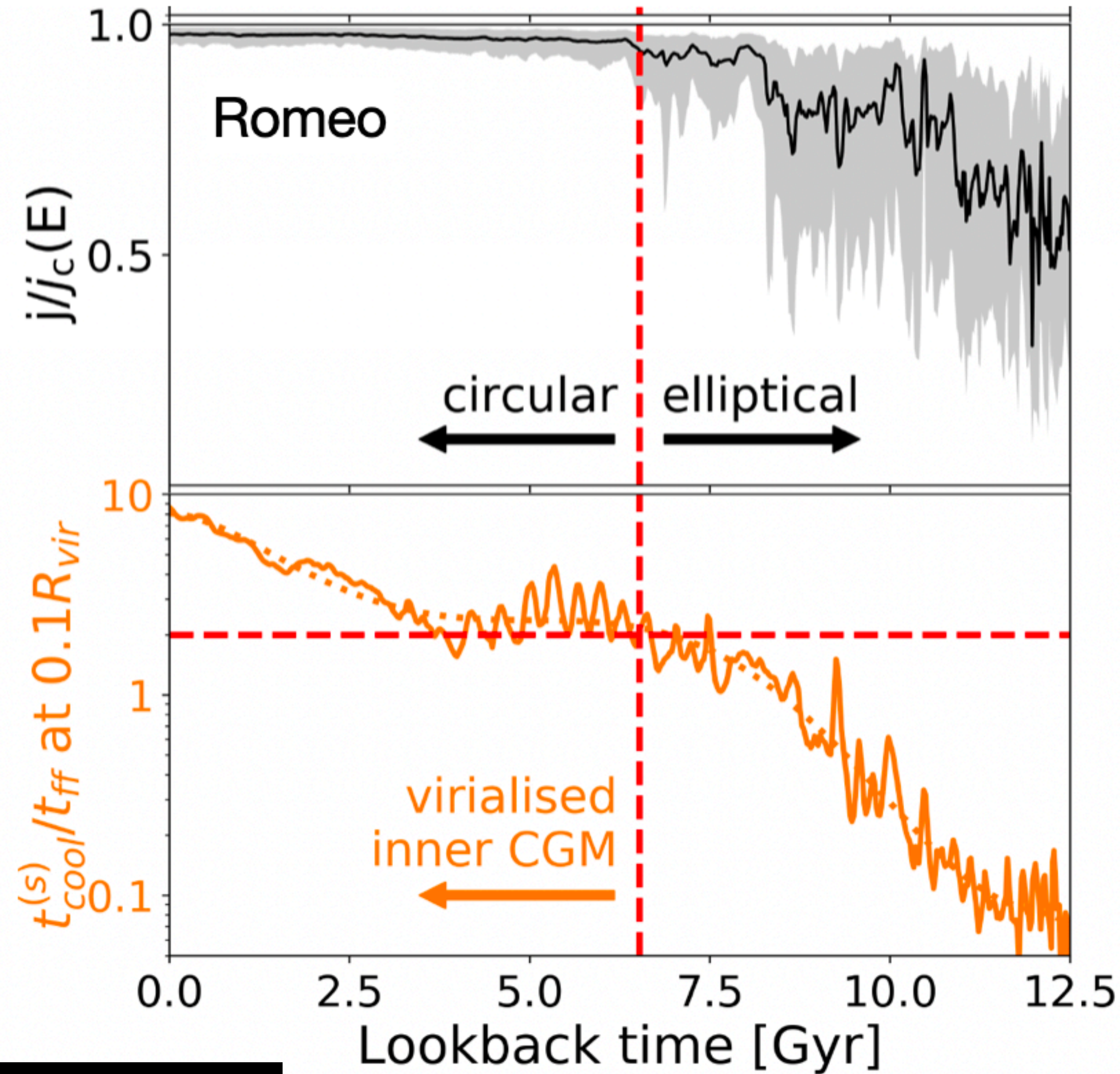
z = 1 \longrightarrow **z = 0**



**Cool inflows crash
supersonically**



**Gentle, coherent
subsonic flows**



*** Young-star orbits become VERY circular after inner CGM virializes**

Why? Hot (sub-sonic) accretion allows angular momentum to mix/align prior to entering the galaxy

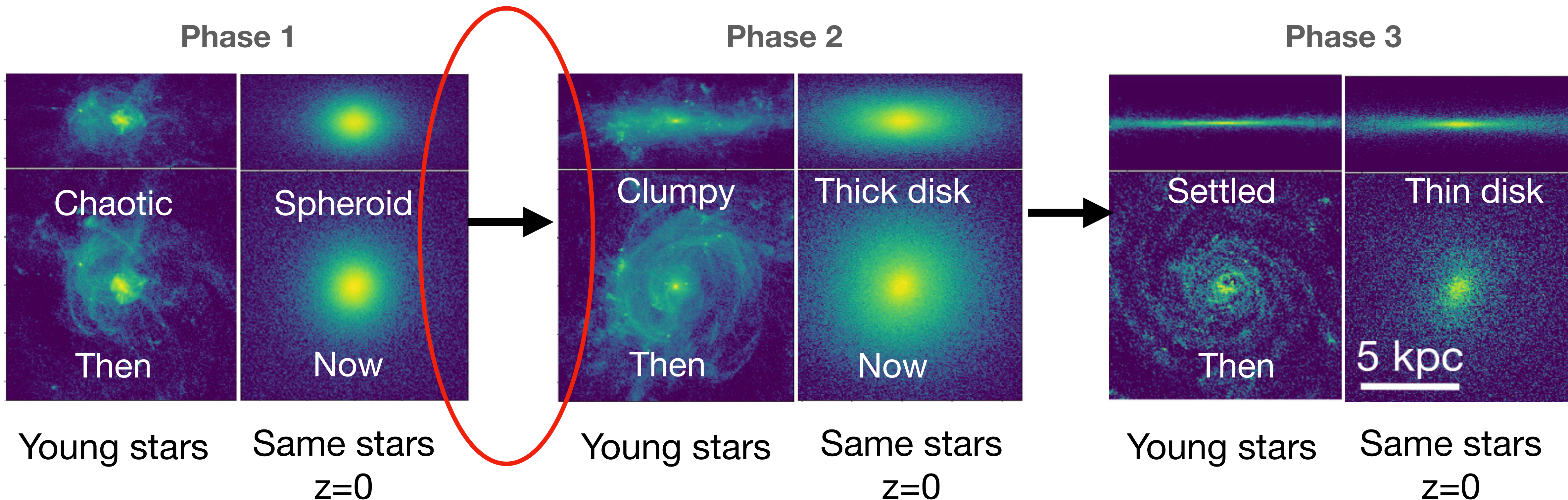
Hafen et al. 2022

Yu+2023

3 Phases of Galaxy Disk Formation

Yu+2023

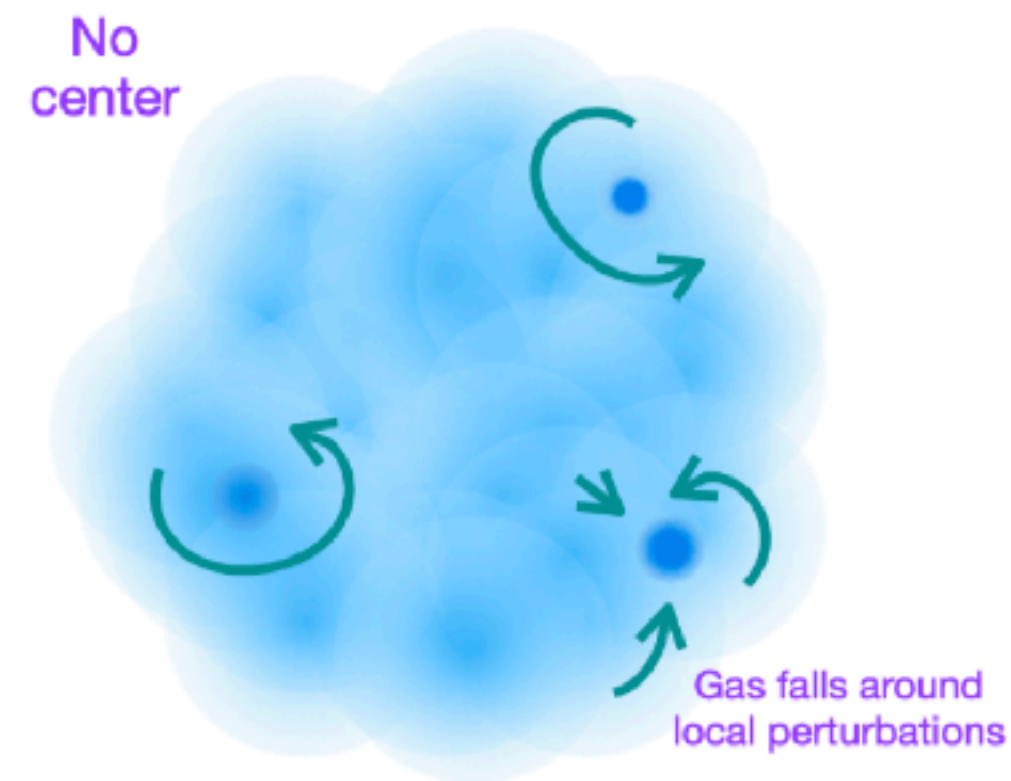
?????



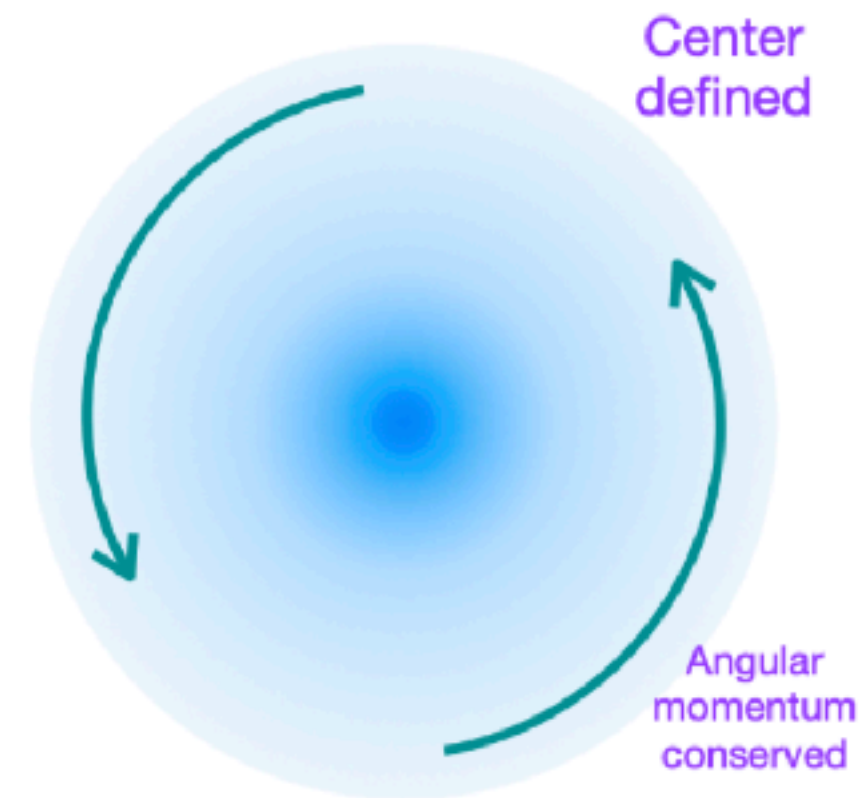
What causes 1 \rightarrow 2 transition?

Disk formation is promoted when the mass profile becomes sufficiently centrally-concentrated

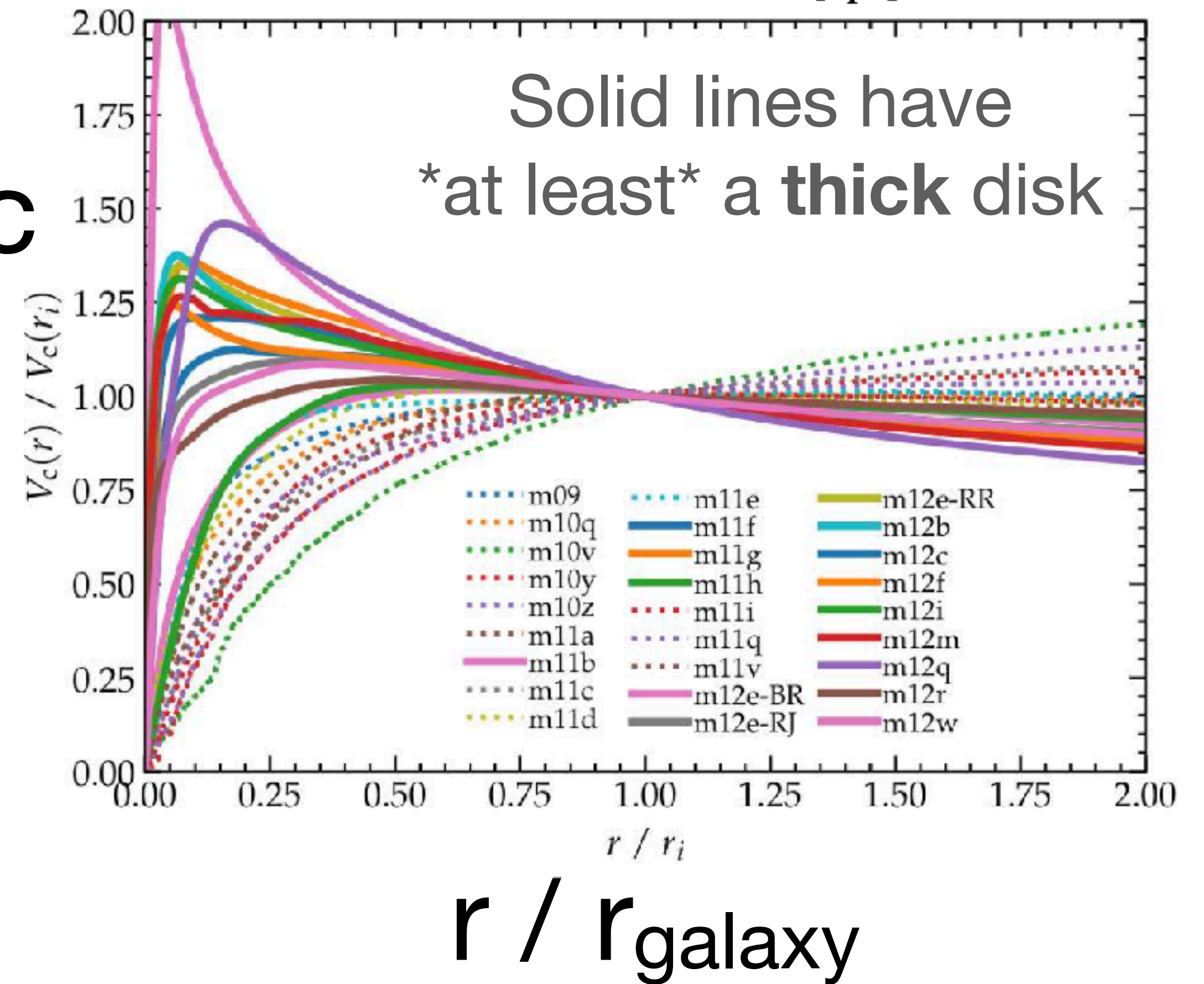
UN-CONCENTRATED PROFILE

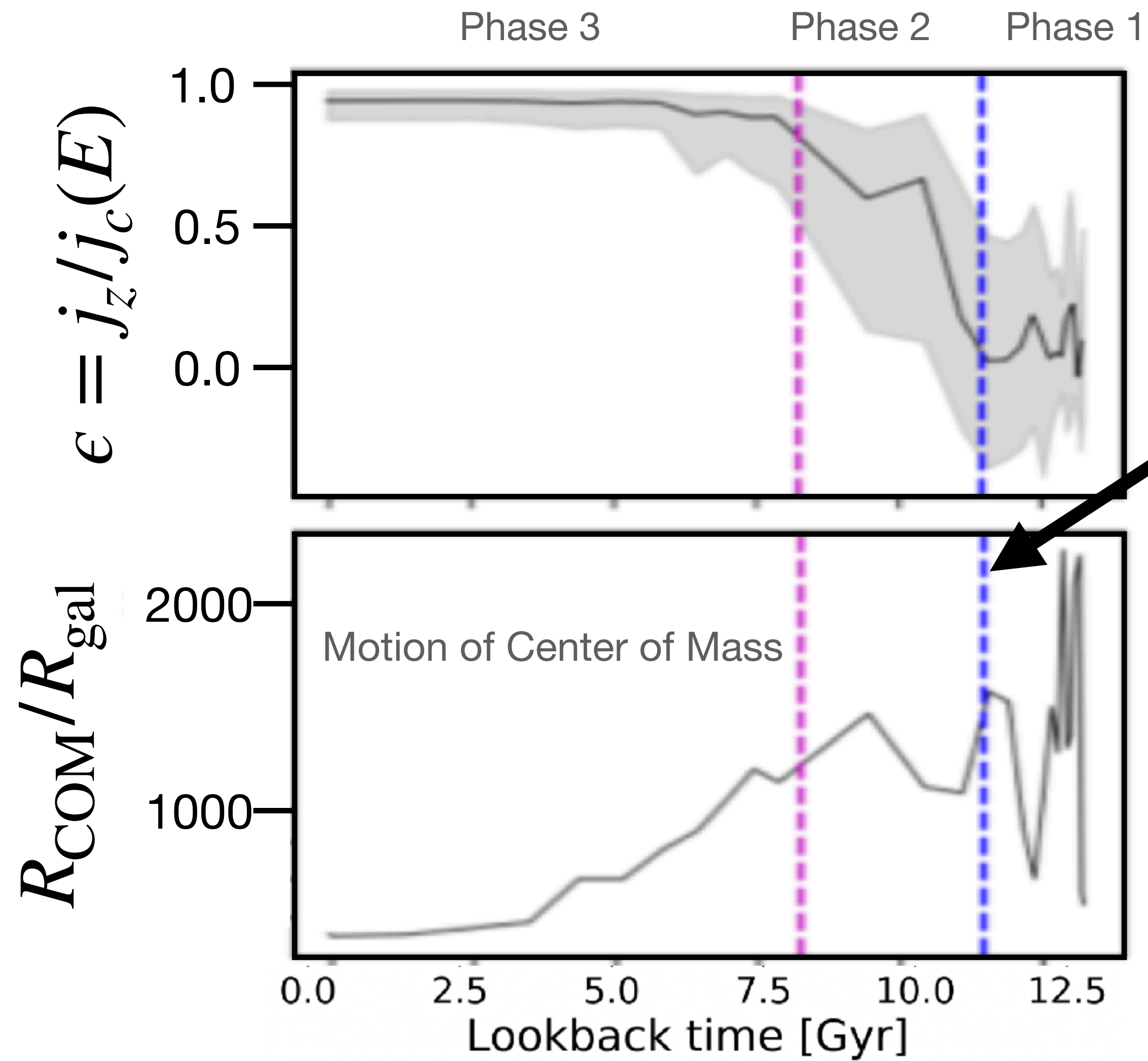


CONCENTRATED PROFILE



V_c

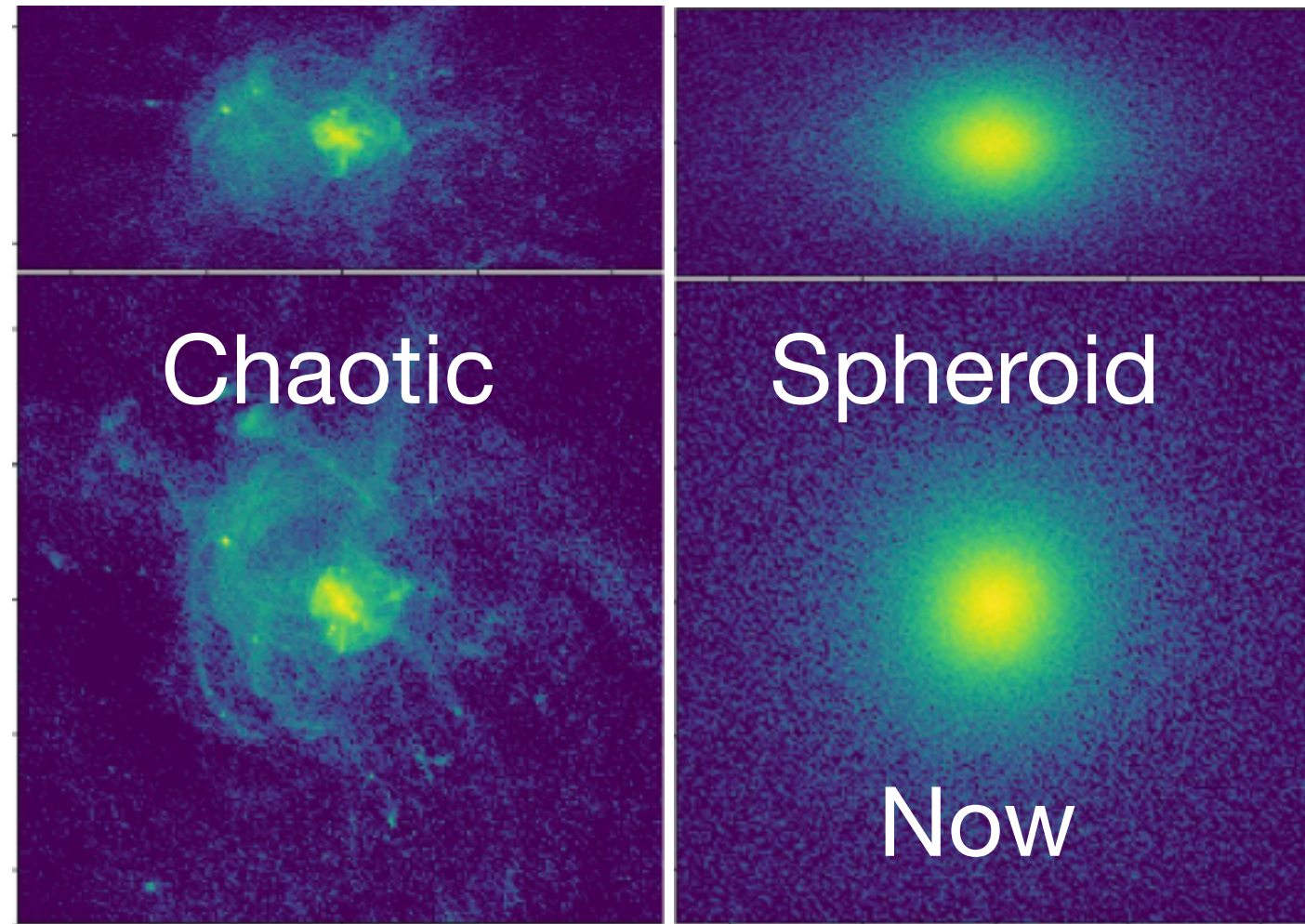




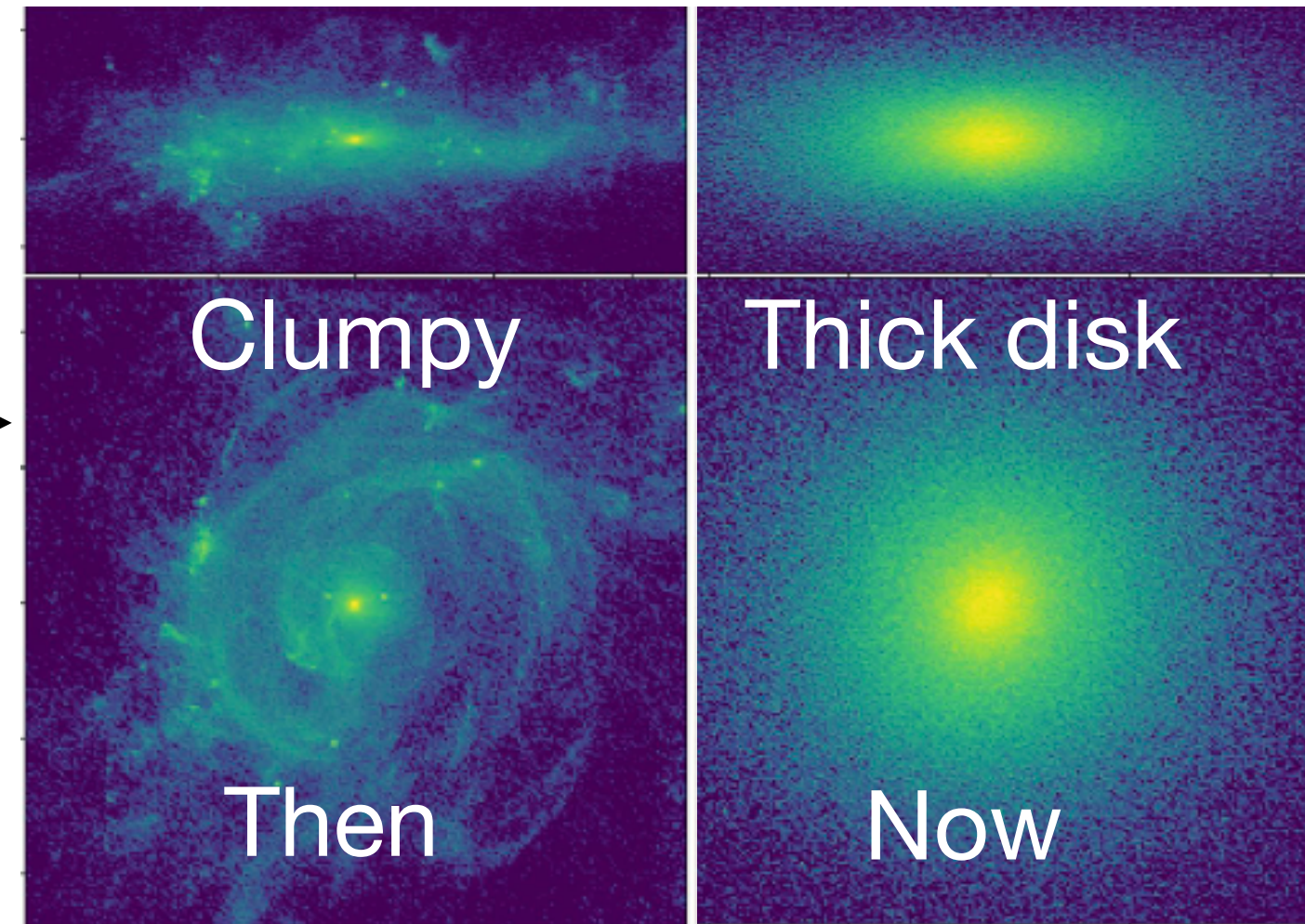
Galaxy “spins up” only after it has a well-established center and concentrated potential.

Olti Myrtaj et al., in prep.

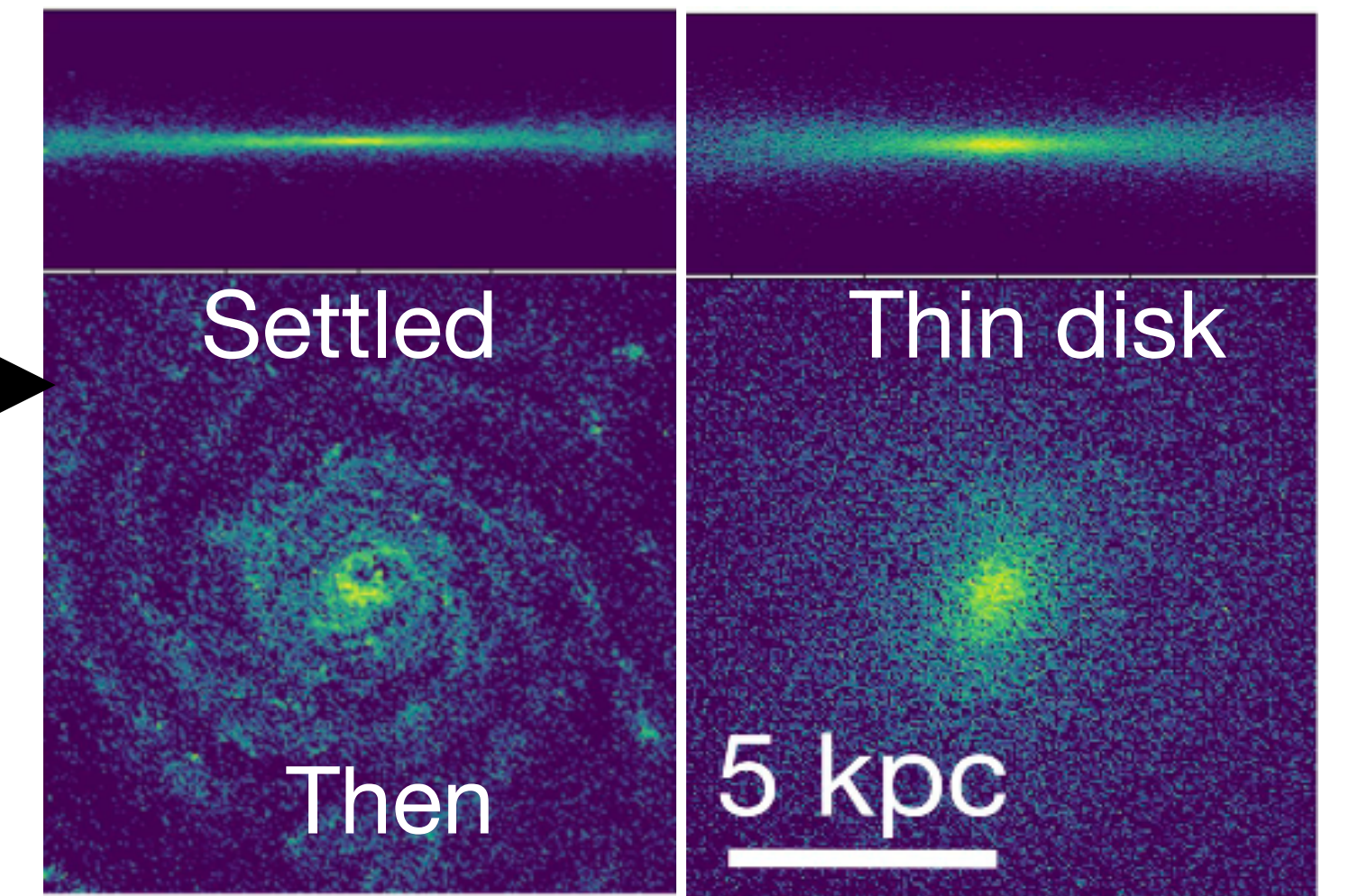
Phase 1



Phase 2



Phase 3



Center of mass is stable; Potential becomes concentrated?

Inner CGM virializes

3 phases of disk formation in FIRE-2

- Phase 1: Chaotic morphology, bursty star formation => spheroid orbits today
- Phase 2: Clumpy, puffy disk, bursty star formation => thick disk orbits today
- Phase 3: Thin, regular disk, steady star formation => (mostly) thin-disk orbits today
- **Thick disk “spin-up” phase enabled by centrally-concentrated potential?**
- **Thin disk “cool down” phase enabled by sub-sonic “hot mode” accretion; angular momentum coherence**

